Optimization of the KAGRA sensitivity

Yuta Michimura

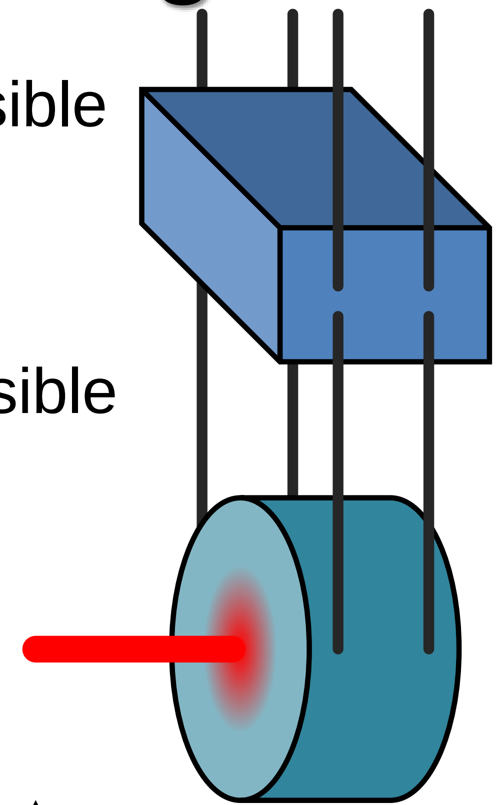
Department of Physics, University of Tokyo

Kentaro Komori, Atsushi Nishizawa, Hiroki Takeda,
Koji Nagano, Yutaro Enomoto, Kazuhiro Hayama,
Kentaro Somiya, Masaki Ando, Sadakazu Haino

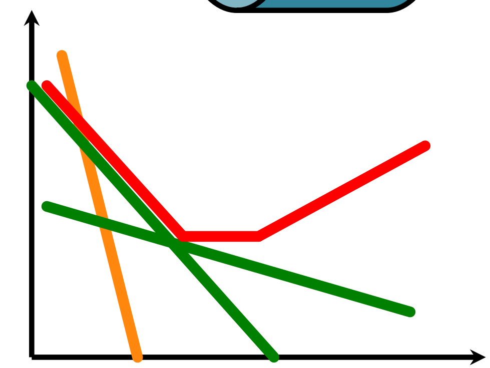
Overview

- Developed a new way to optimize the KAGRA sensitivity design based on
  - CBC inspiral range
  - CBC parameter estimation
- Optimization done by Particle Swarm Optimizer
  YM+, [Phys. Rev. D 97, 122003 (2018)]
- Studied possible KAGRA+ candidates with budget constraints
  - 40 kg mirror with better coating
  - 400 W laser with squeezing
  - Frequency dependent squeezing

Room Temperature Design



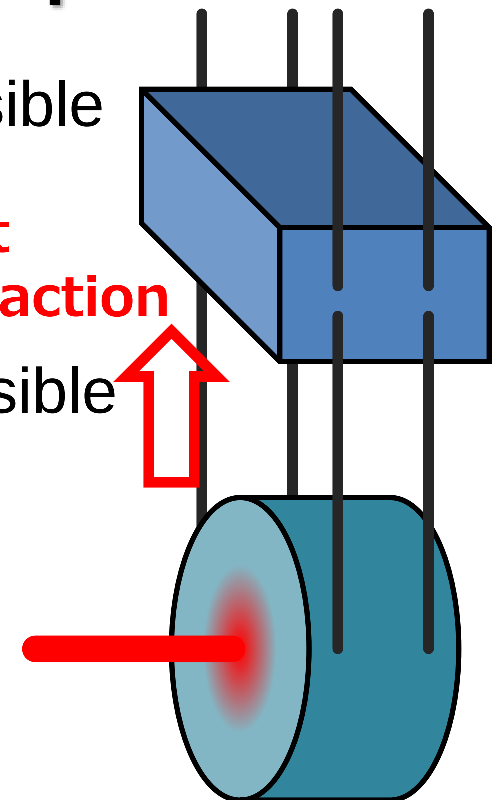
- Seismic noise: reduce as much as possible
  - multi-stage suspensions
  - underground
- Thermal noise: reduce as much as possible
  - larger mirror
  - thinner and longer fibers



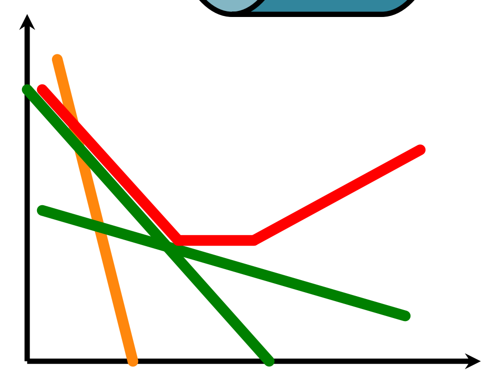
- Quantum noise: optimize the shape
  - input laser power
  - tune signal recycling parameters

Cryogenic Design is Complicated

- Seismic noise: reduce as much as possible
  - multi-stage suspensions
  - underground



- Thermal noise: reduce as much as possible
  - larger mirror
  - thinner and longer fibers
  - cryogenic cooling
- Quantum noise: optimize the shape
  - input laser power
  - tune signal recycling parameters



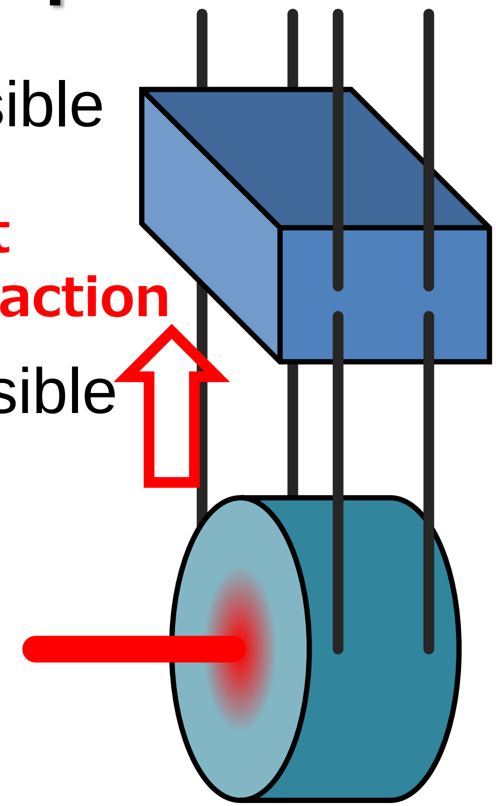
Cryogenic Design is Complicated

- Seismic noise: reduce as much as possible
  - multi-stage suspensions
  - underground
- Thermal noise: ~~reduce as much as possible~~ optimize
  - larger mirror
  - thinner and longer fibers
  - cryogenic cooling

worse cooling power
mirror heating

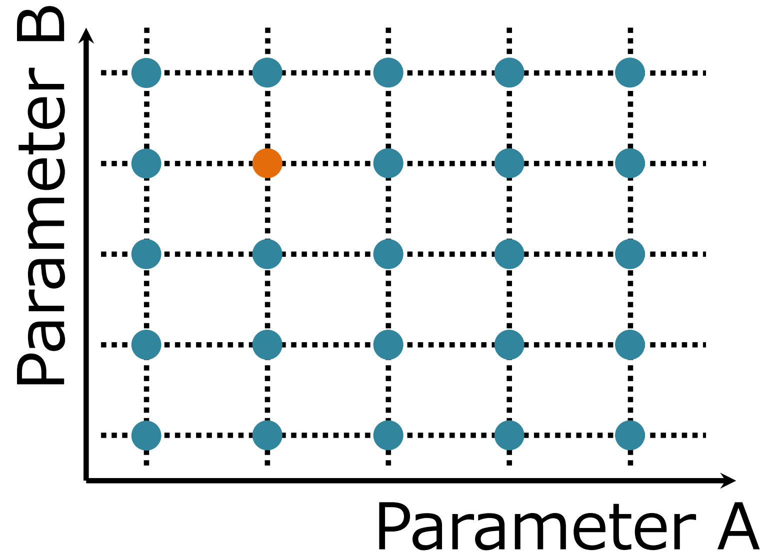
- Quantum noise: optimize the shape
  - input laser power
  - tune signal recycling parameters

heat extraction



Grid-based Search is not Scalable

- Sensitivity design is an optimization problem
- Grid-based parameter search
  - deterministic
  - computational cost grows exponentially with number of parameters



- Future GW detectors (like KAGRA+) require more parameters to be optimized
- Almost impossible with grid-based approach

Particle Swarm Optimization!

- Particles search the parameter space based on own best position and entire swarm's best known position

$$x_k(t+1) = x_k(t) + v_k(t)$$

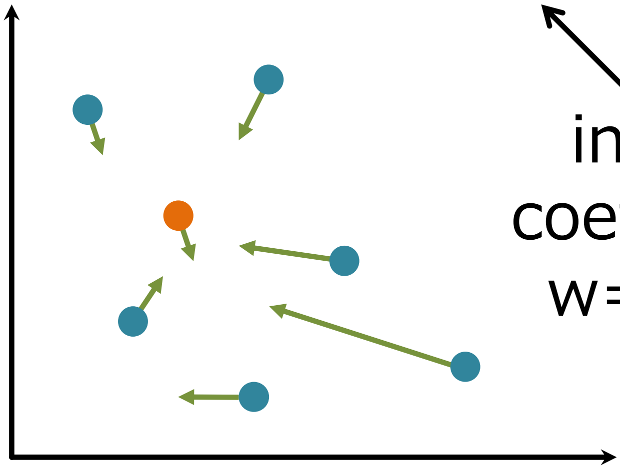
personal best position so far

global best position so far

$$v_k(t+1) = wv_k(t) + c_1 r_1(\hat{x}_k - x_k(t)) + c_2 r_2(\hat{x}_g - x_k(t))$$

inertia coefficient w=0.72

acceleration coefficient c=1.19
random number r ∊ [0,1]

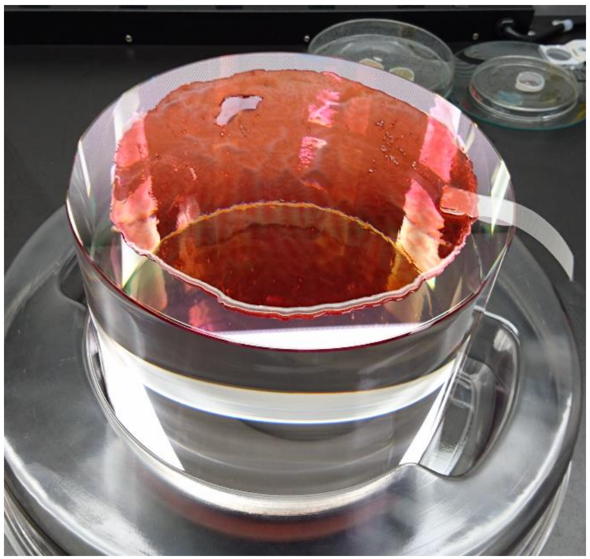


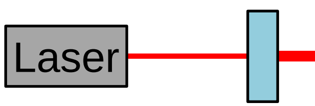
Parameter space

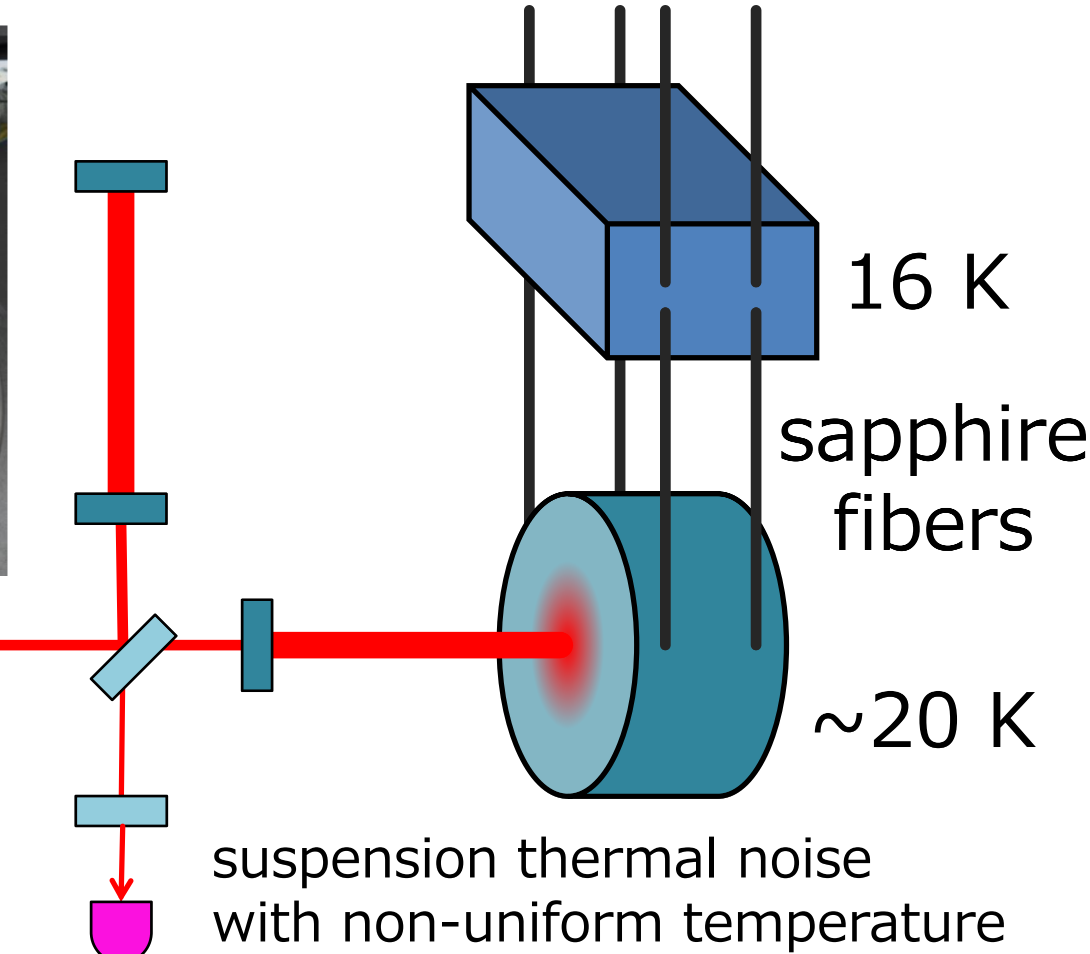
Kennedy & Eberhart (1995)
values for w and c are from Standard PSO 2006

Apply PSO for KAGRA Design

- RSE interferometer
- Cryogenic sapphire test masses



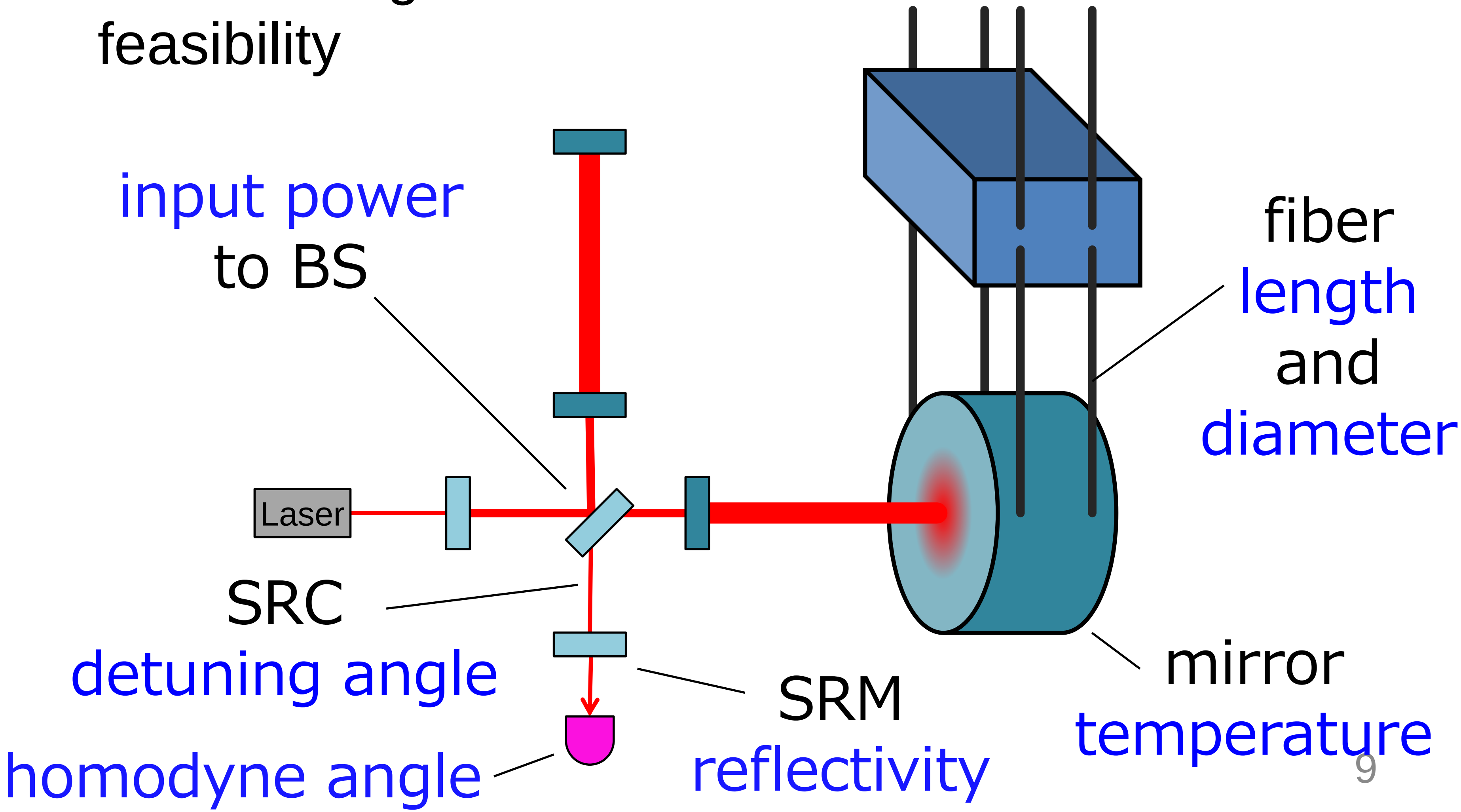




suspension thermal noise
with non-uniform temperature
K. Komori+ [PRD 97 102001 (2018)]

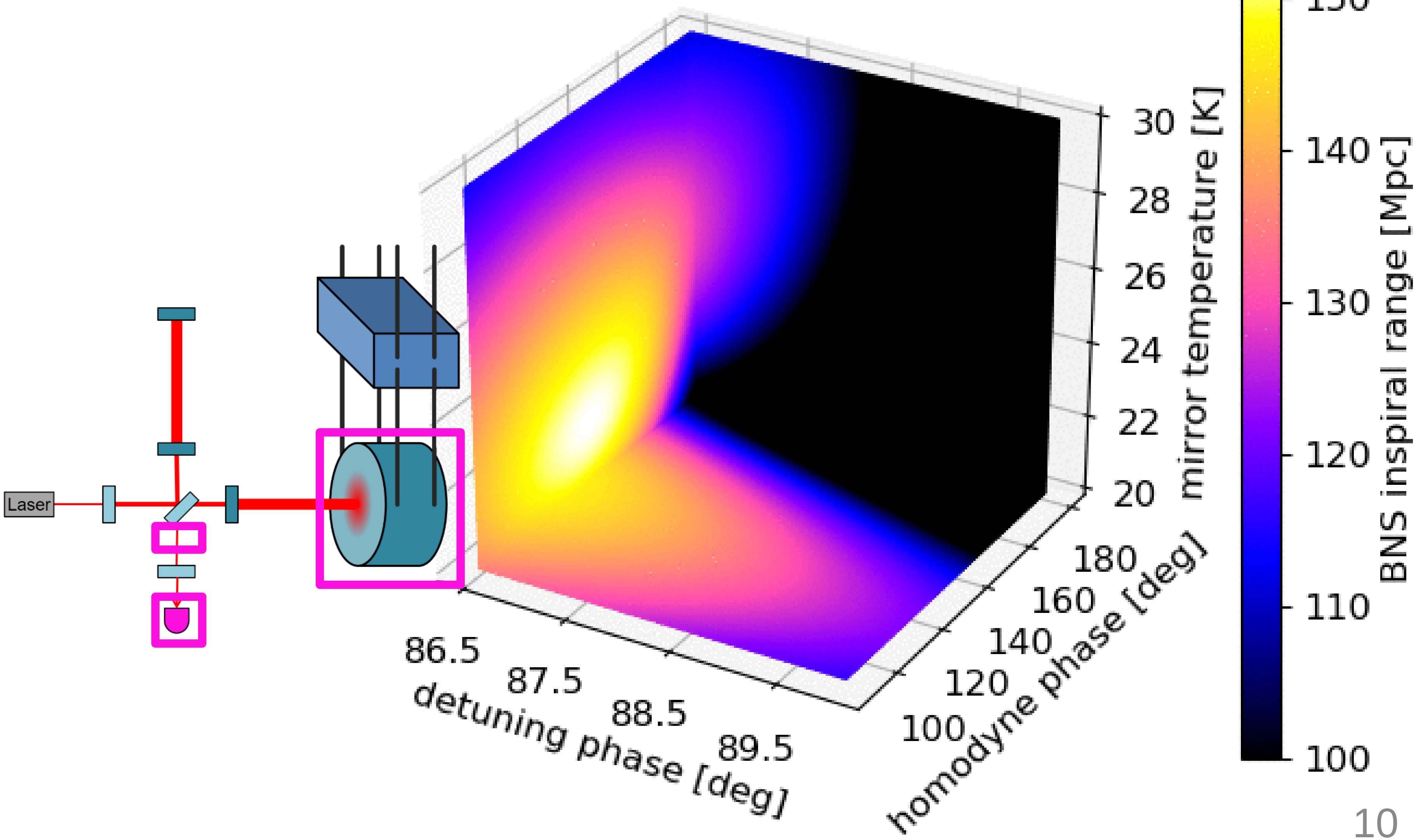
Parameters of Interest

- 7 parameters are relatively easy to be retuned
- Search range based on feasibility



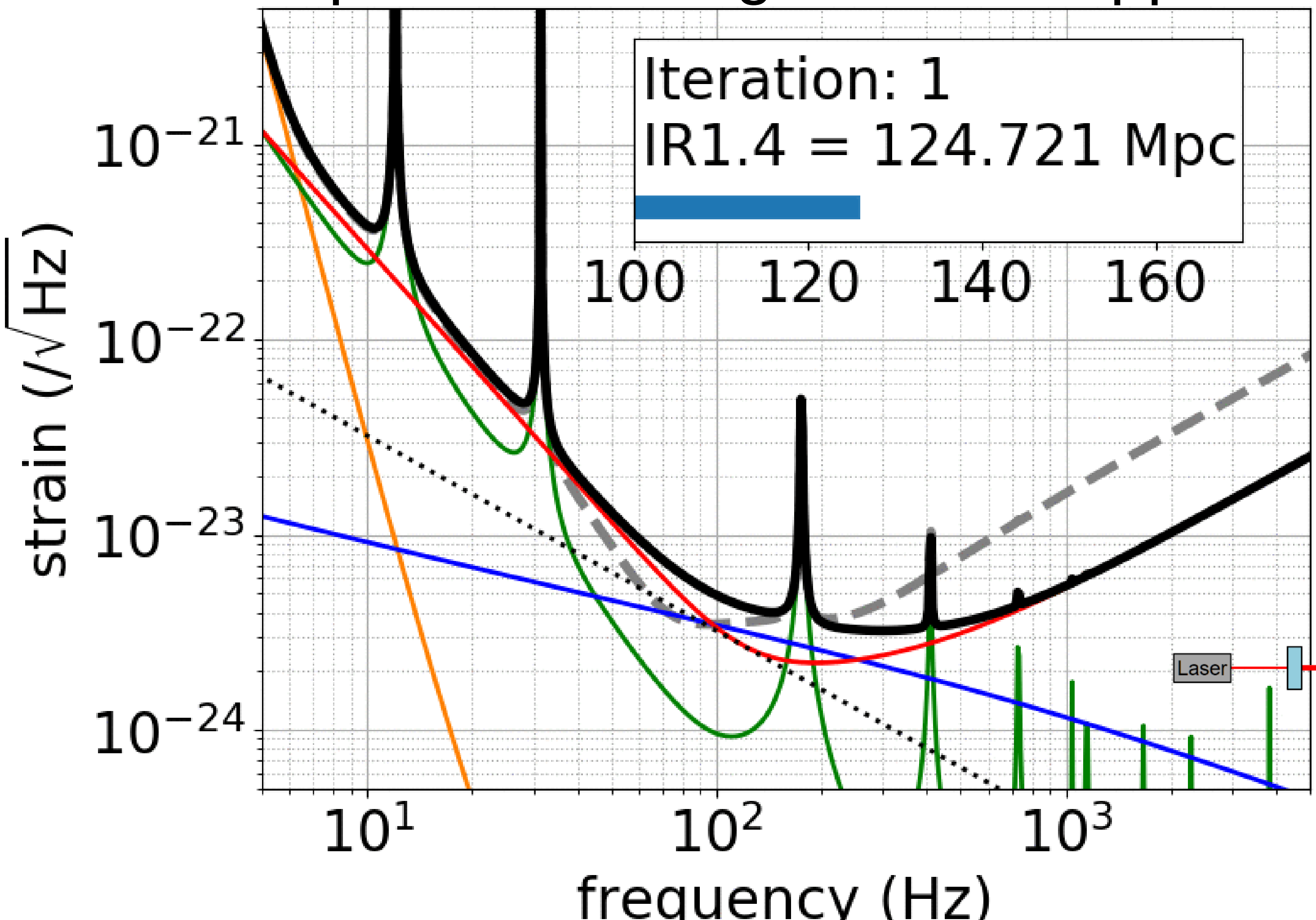
3 Parameter Optimization

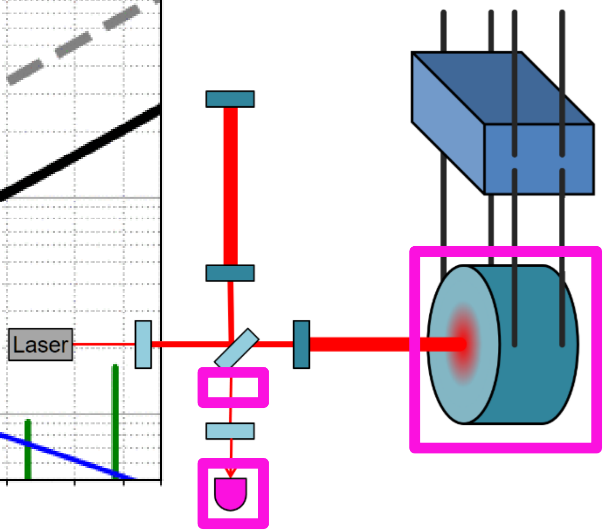
- Consistent with grid-based approach



3 Parameter Optimization

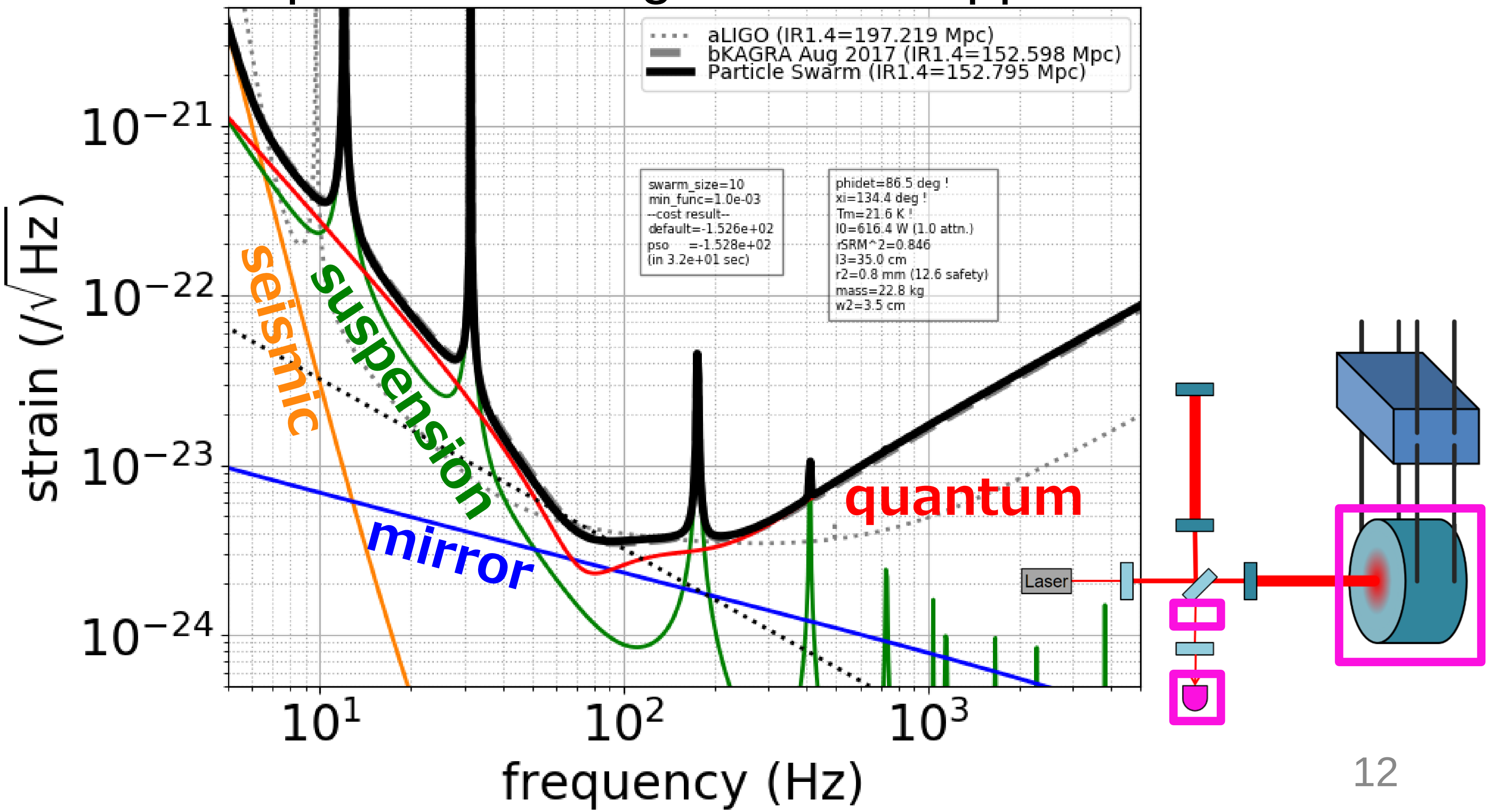
- Consistent with current designed sensitivity which was optimized with grid-based approach





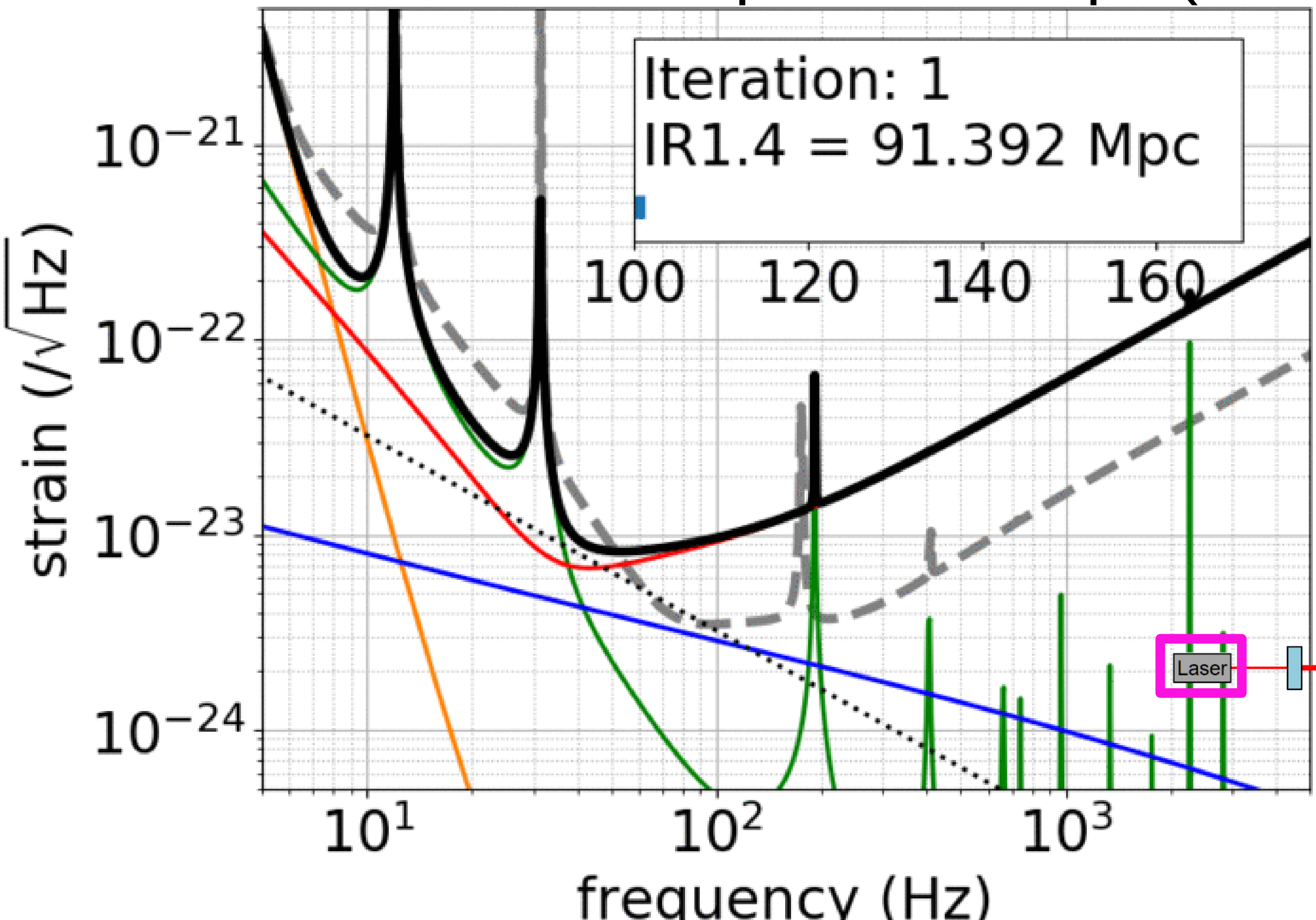
3 Parameter Optimization

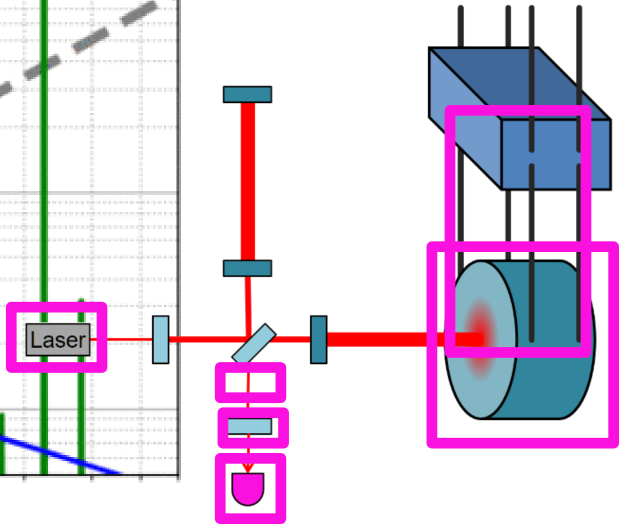
- Consistent with current designed sensitivity which was optimized with grid-based approach



7 Parameter Optimization

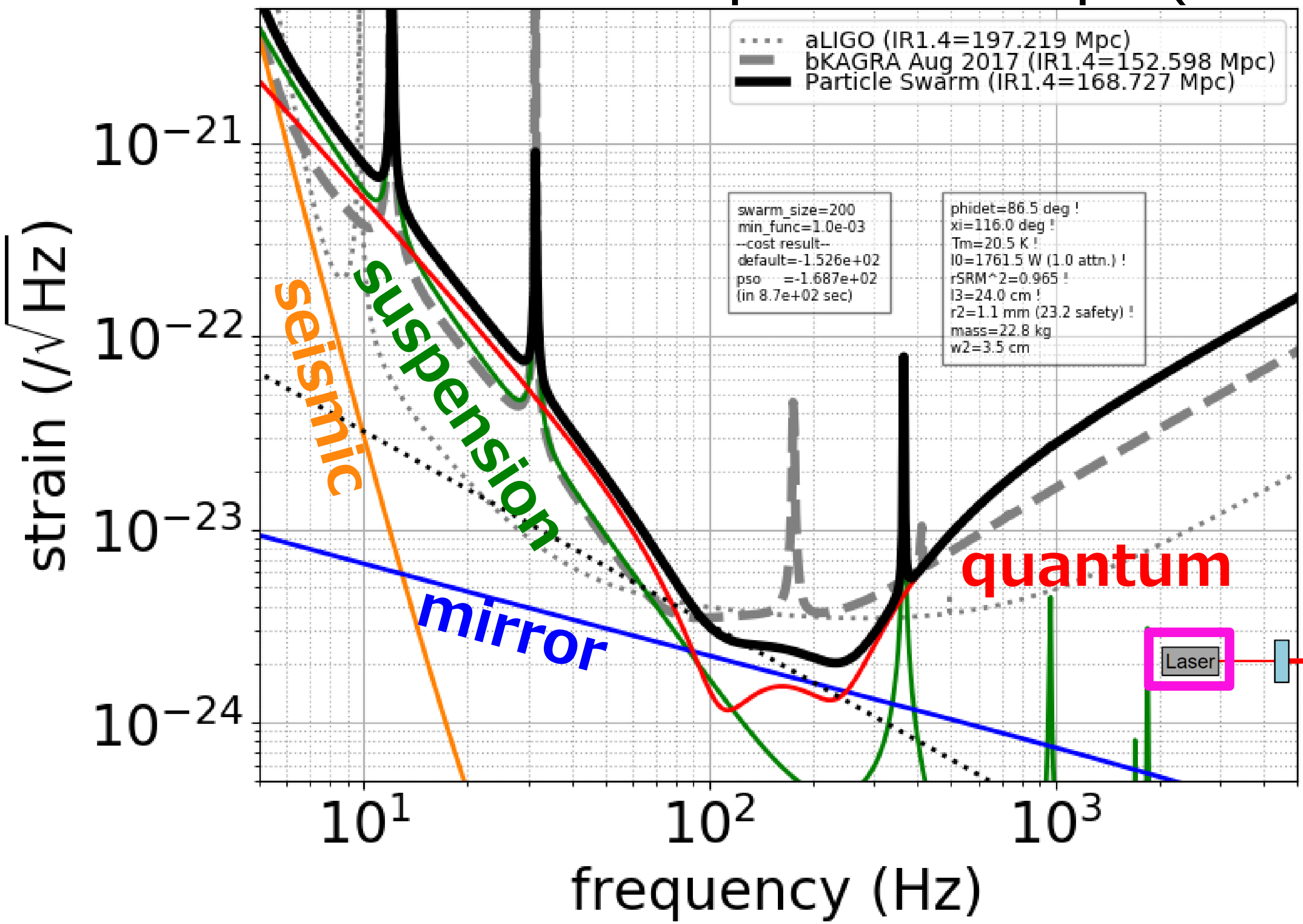
- Changing suspension fibers and SRM increases BNS IR from 153 Mpc to 169 Mpc (10% increase)



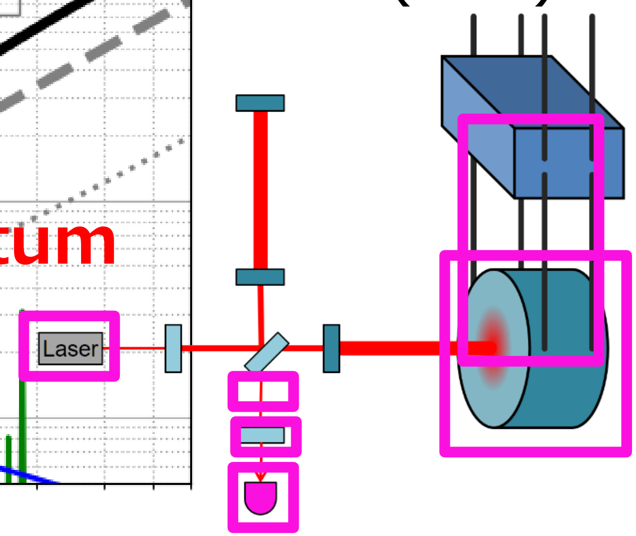


7 Parameter Optimization

- Changing suspension fibers and SRM increases BNS IR from 153 Mpc to 169 Mpc (10% increase)



shorter and thicker to put more power with 20 K (min)



Sensitivity Design with PSO is Fast

- Optimization done in O(10) minutes with my laptop
- Number of cost function evaluations

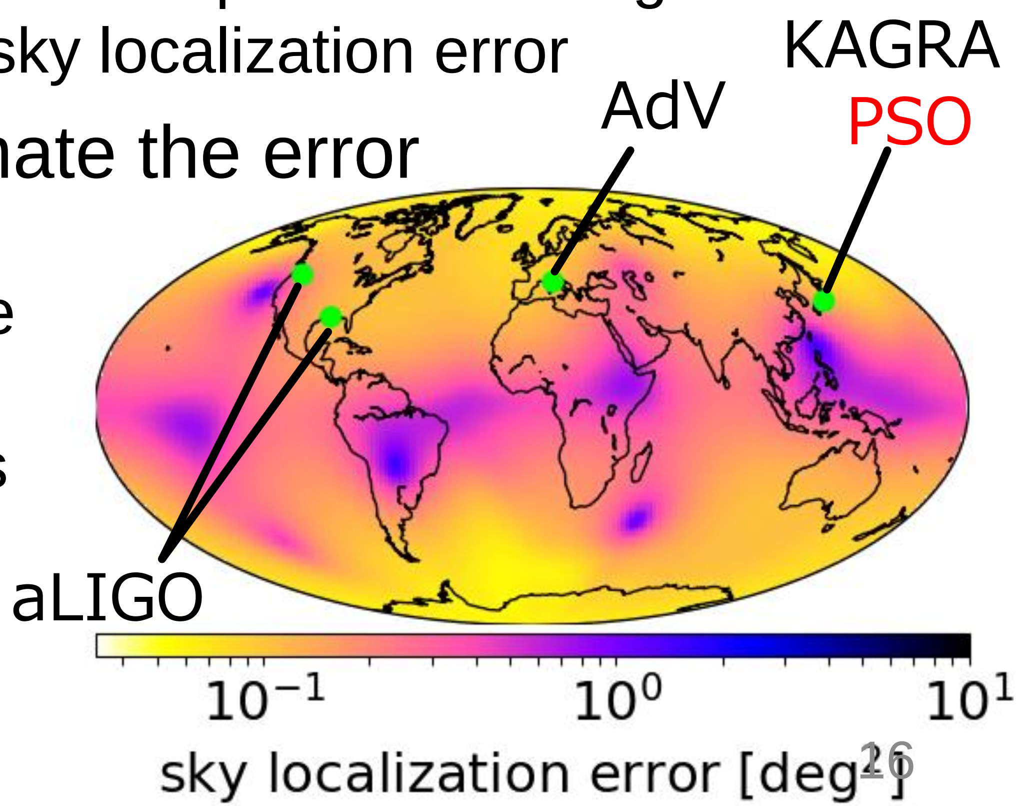
	Grid-based	PSO
3 params	$\sim 10^5$	$10 \times (52 \pm 13)$
5 params	$\sim 10^9$	$20 \times (73 \pm 16)$
7 params	$\sim 10^{14}$	$200 \times (60 \pm 18)$

* In case optimization is done at precision of 0.1 Mpc

- Computational cost do not grow exponentially with dimensionality of parameter space
- Useful for optimization with many parameters, computationally expensive cost function

Sky Localization Optimization

- Cost function:
  sky localization of GW170817-like binary
  - 1.25-1.5 Msun at 40 Mpc, inclination 28 deg
  - zero spins, no precession
  - 108 sets of sky location and polarization angle to derive median of sky localization error
- Fisher matrix to estimate the error
  - inspiral waveform to
    3.0 PN in amplitude
    3.5 PN in phase
  - 11 binary parameters
- HLVK global network

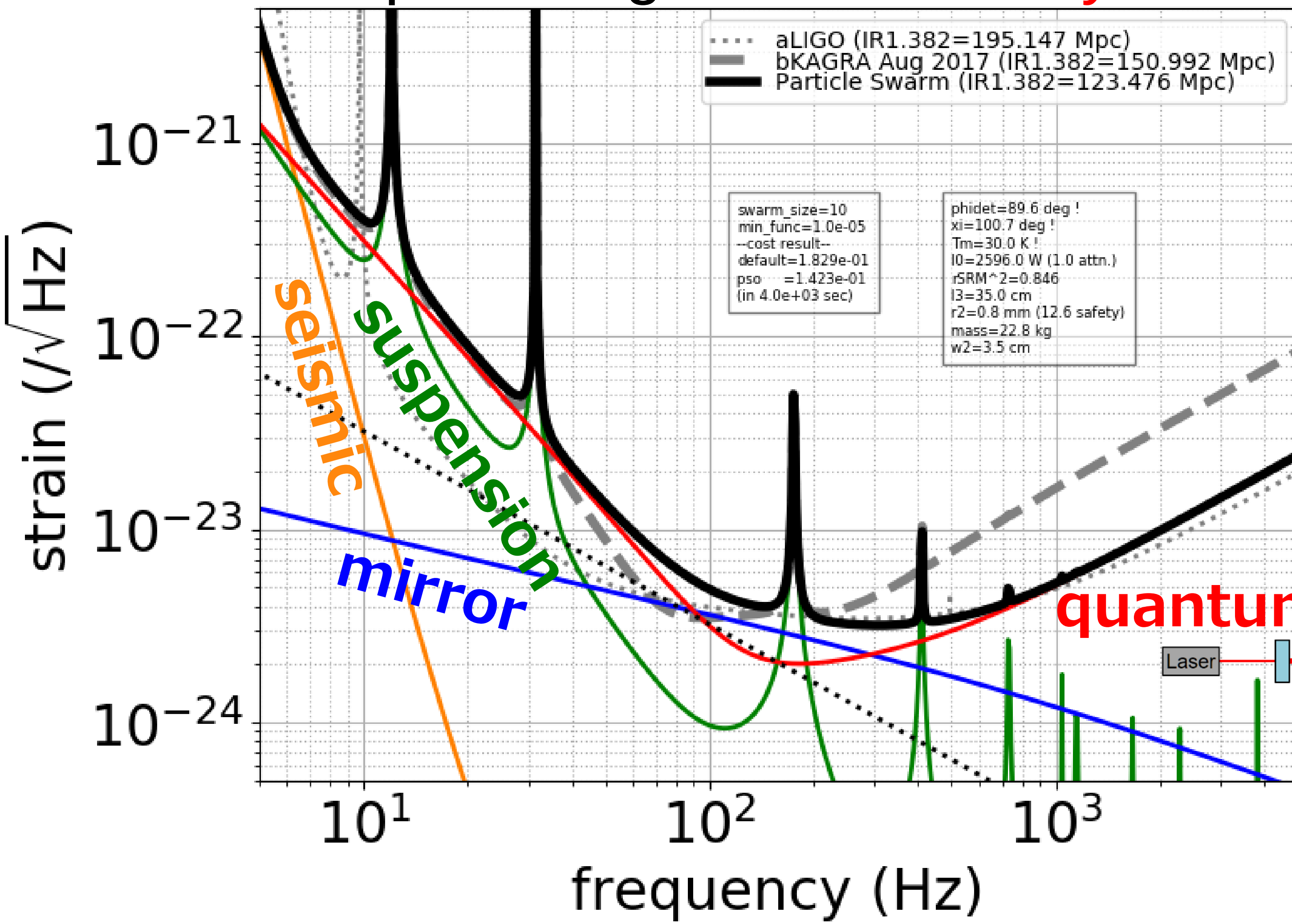


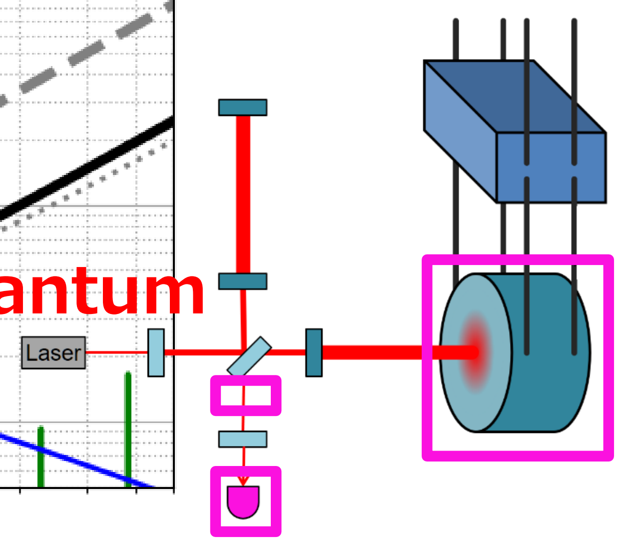
3 Parameter Optimization

- Sky localization improved by a factor of 1.3 but inspiral range is reduced by 20%

0.18 deg^2 → 0.14 deg^2

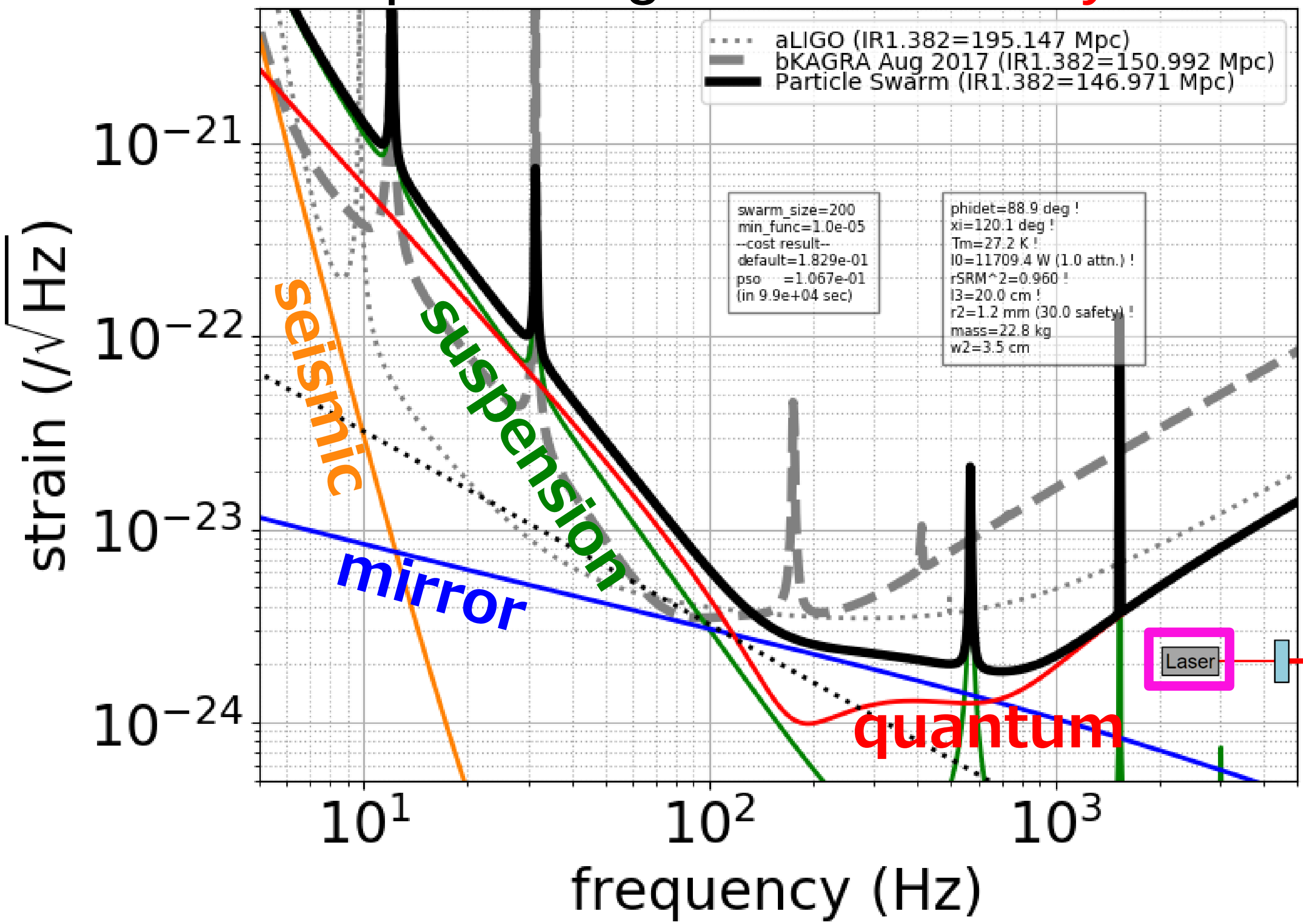
2600 W at BS
30 K





7 Parameter Optimization

- Sky localization improved by a factor of 1.6 but inspiral range is reduced by 2%

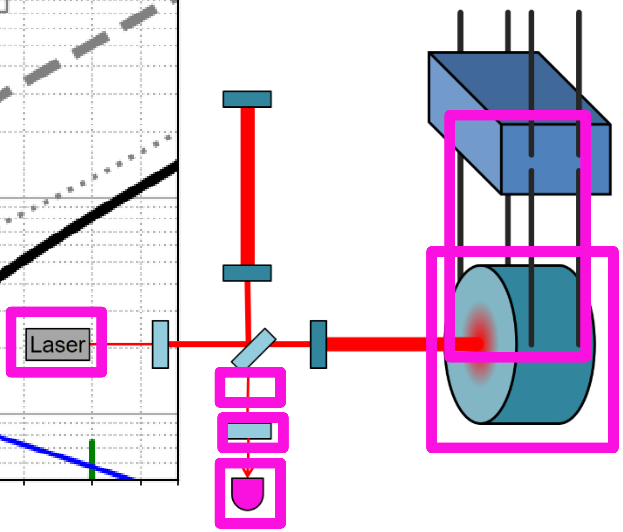


$0.18\ \text{deg}^2 \rightarrow 0.11\ \text{deg}^2$

shortest and thickest fibers

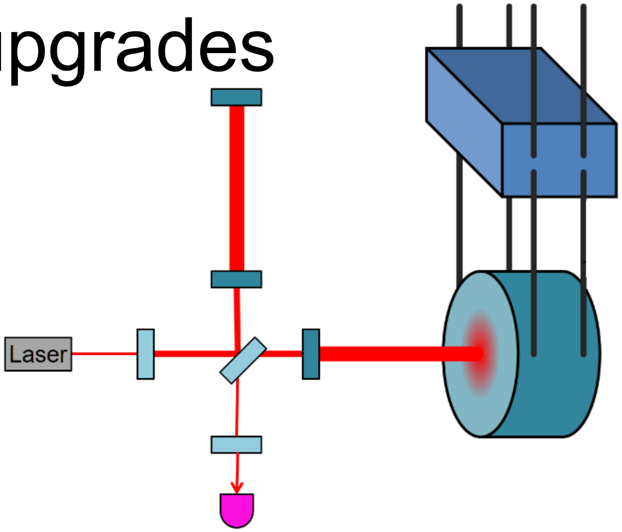
12 kW at BS

27 K



KAGRA+ with Budget Constraints

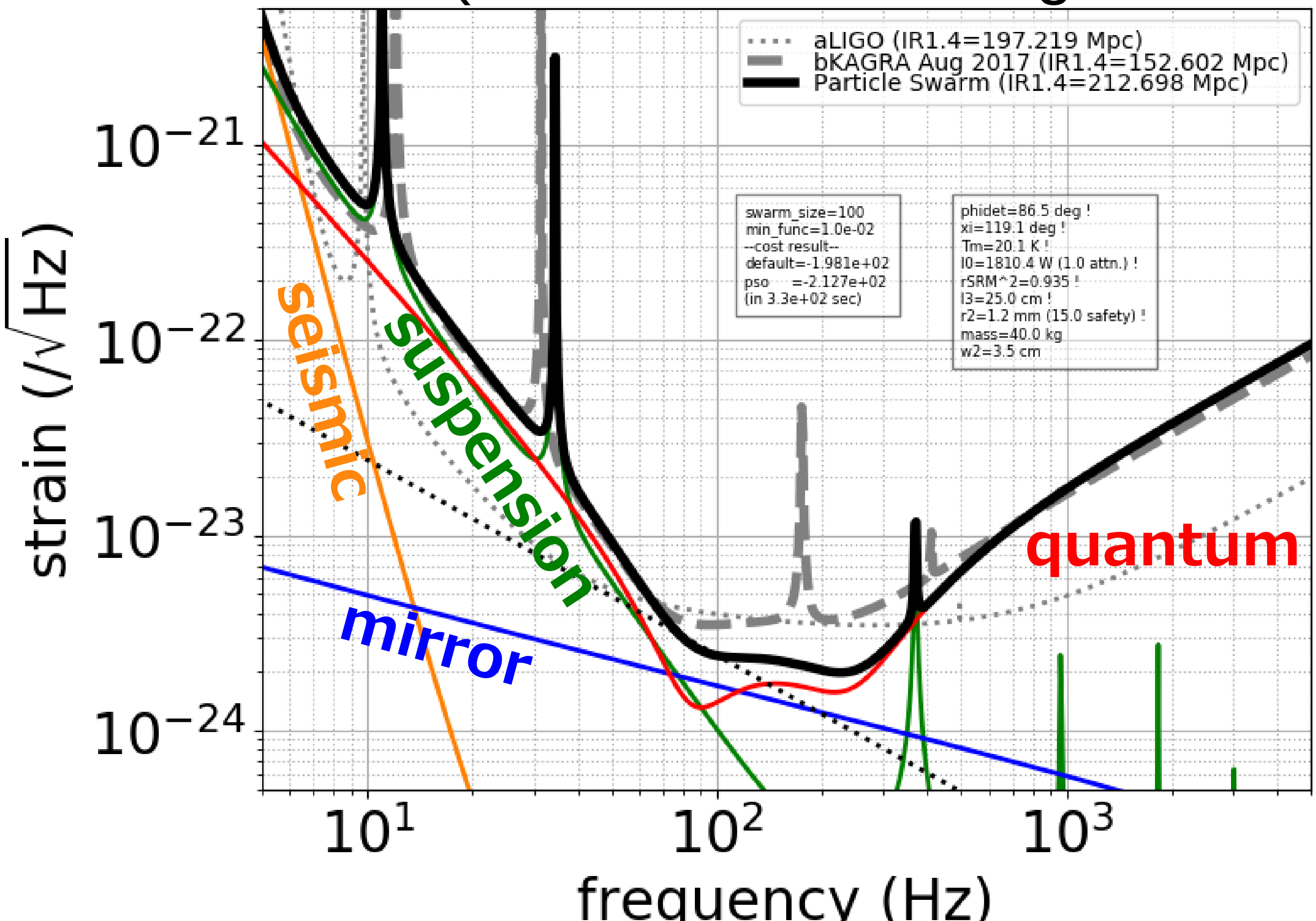
- Let’s consider a bit more drastic upgrades
- Suppose you have $5M for KAGRA+



- Candidates would be
  A. 40 kg mirror with better coating (>$4M?)
  and new sapphire fibers ($1M?)
  (use existing cryostat and Type-A tower)
  B. 400 W laser ($3M?) with squeezing ($1M?)
  and new sapphire fibers ($1M?)
  C. Frequency dependent squeezing ($3M?)
  and new sapphire fibers ($1M?)

Plan A: 40 kg Mirror

- Also assumes factor of 2 coating loss angle reduction (no beam size change assumed)

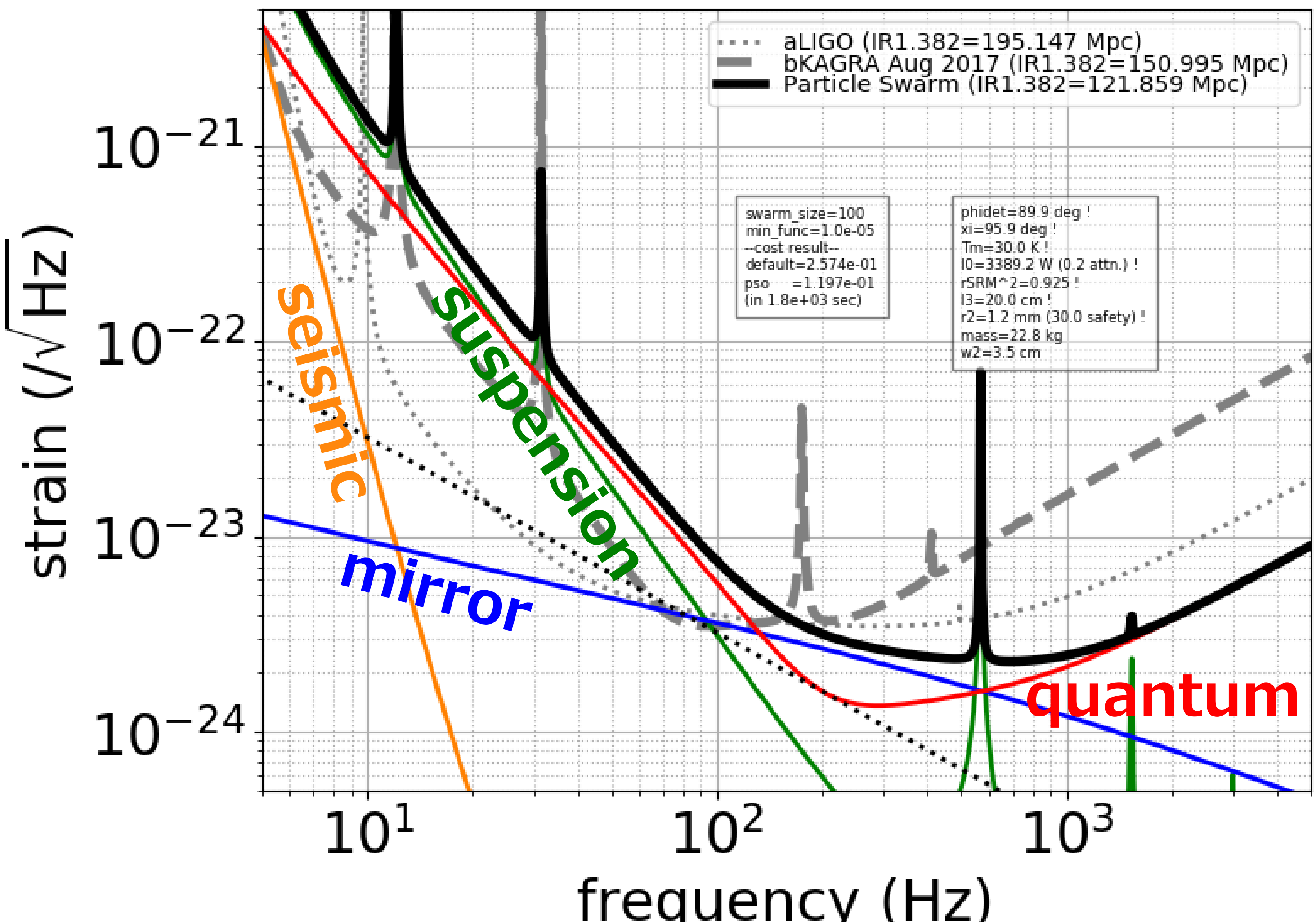


Good for mid frequency improvement → BNS range optimized

T=20.1 K
181 W input
thicker fiber
25.0 cm
φ1.2 mm
(thicker to allow for higher power)

Plan B: 400 W Laser with SQZ

- Assumes 10dB input SQZ

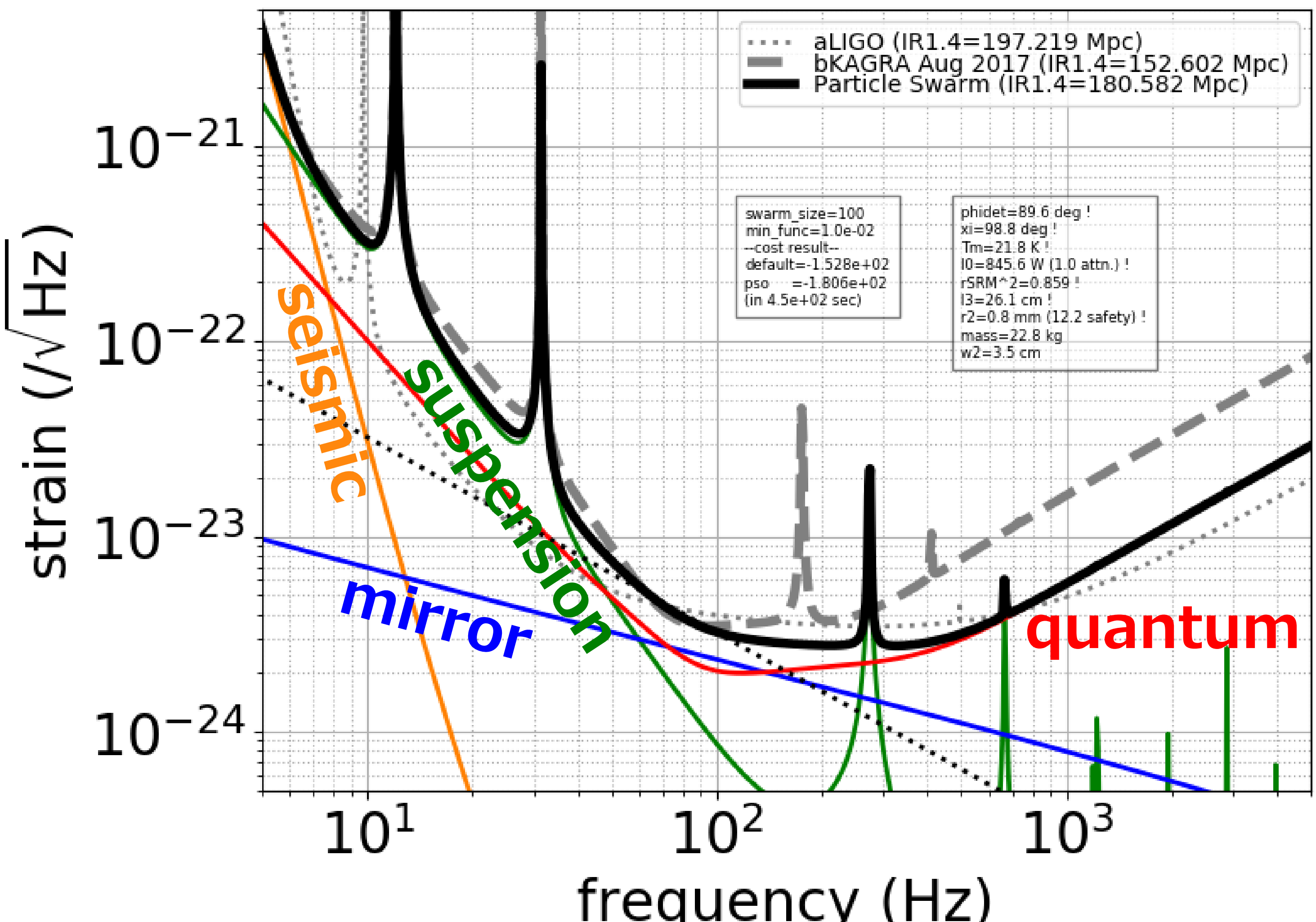


Good for high frequency improvement → BNS range optimized

T=29.8 K
330 W input
shorter and thicker fiber
20.1 cm
φ1.2 mm
(high power with high temperature)

Plan C: Freq. Dependent SQZ

- Assumes 10dB input SQZ and 100 m filter cavity



Broadband improvement → BNS range optimized

T=21.8 K
85 W input
thinner fiber
 26.1 cm
 φ0.8 mm
(SQZ helps to ease fiber requirement)

Summary of $5M Plans

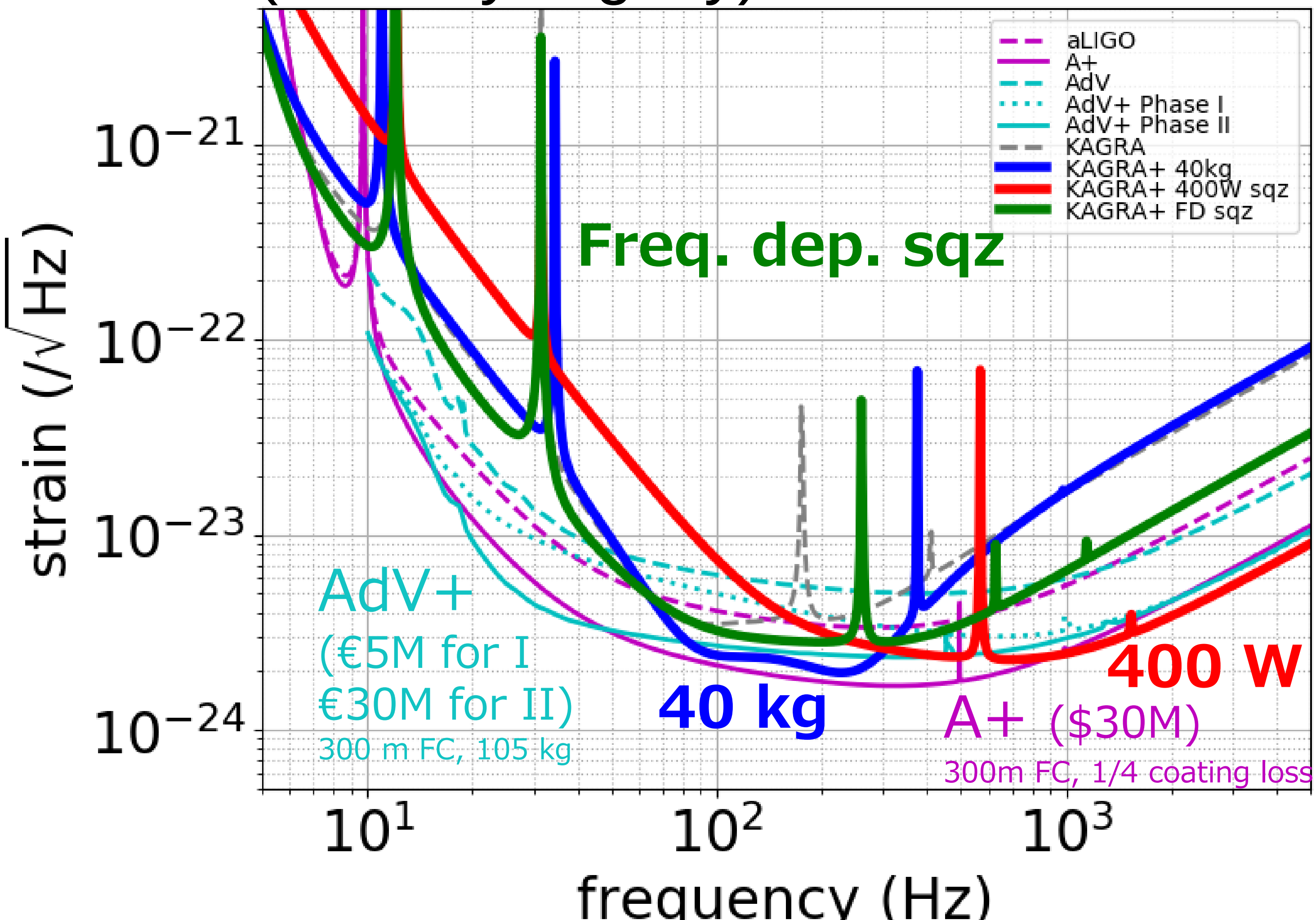
- A. New mirror takes time to fabricate
- B. High power operation is tough
- C. Does it fit in the facility?

	Inspiral range (Mpc)			BNS localize (deg^2)
	BBH100	BBH30	BNS	
bKAGRA	353	1095	153	0.183
A. 40 kg mirror	339	1096	**213**	0.151
B. 400 W laser sqz	117	314	123	**0.114**
C. Freq. dep. sqz	**470**	**1177**	181	0.135

- I like A because of simplicity, but if fabrication of heavier mirrors cannot be done on time, go for C?

Comparison Between 2G and 2G+

- Only Plan B (400W laser with squeezing) can beat A+ (but only slightly)



aLIGO curve from LIGO-T1800044 (updated ver)

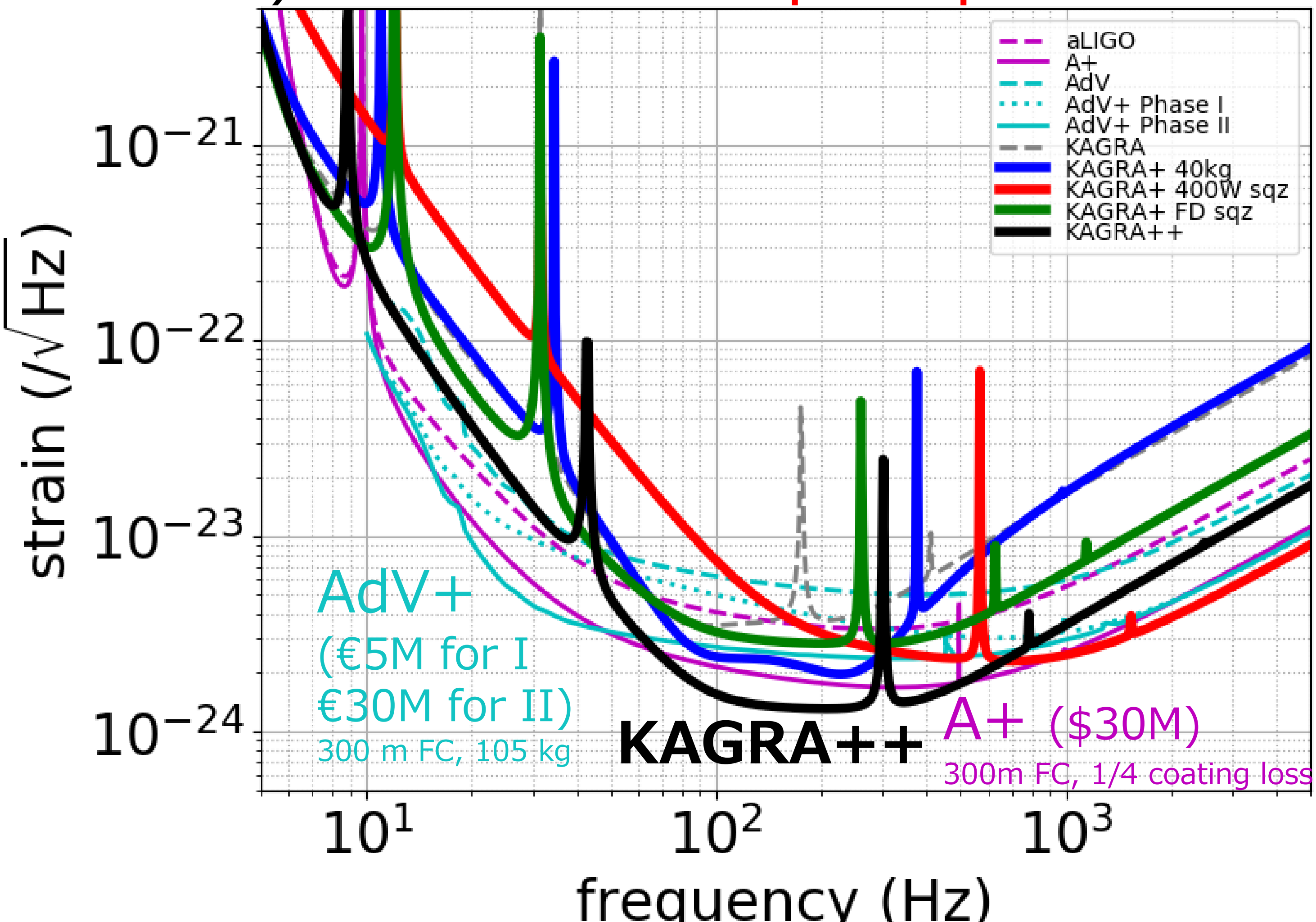
A+ curve from LIGO-T1800042

AdV and AdV+ curves from VIR-0325B-18

KAGRA curve from JGW-T1707038

Be Optimistic, Combine Them!

- 100 kg mirror with1/4 coating loss (and larger beam size), 320 W, 10dB input sqz with 100 m filter cavity



355 Mpc

~ 10 yrs ?

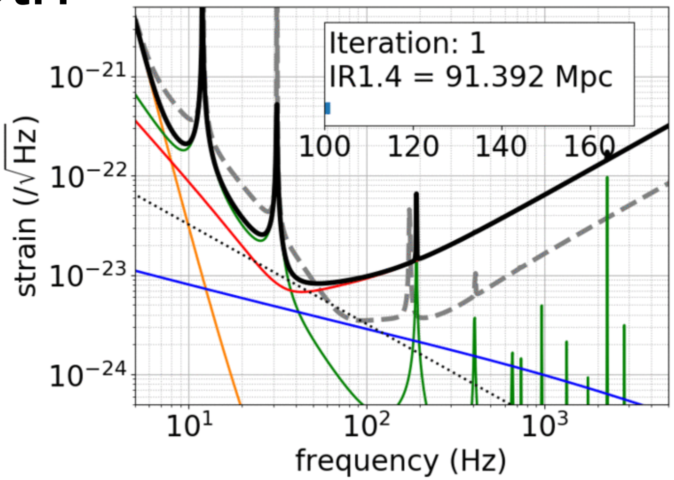
~$20M ?

Summary

- Demonstrated sensitivity design with PSO
- Application to KAGRA shows both
  - BNS inspiral range
  - BNS sky localization

  can be improved by retuning 7 parameters of existing components

  YM+, [Phys. Rev. D 97, 122003 (2018)]



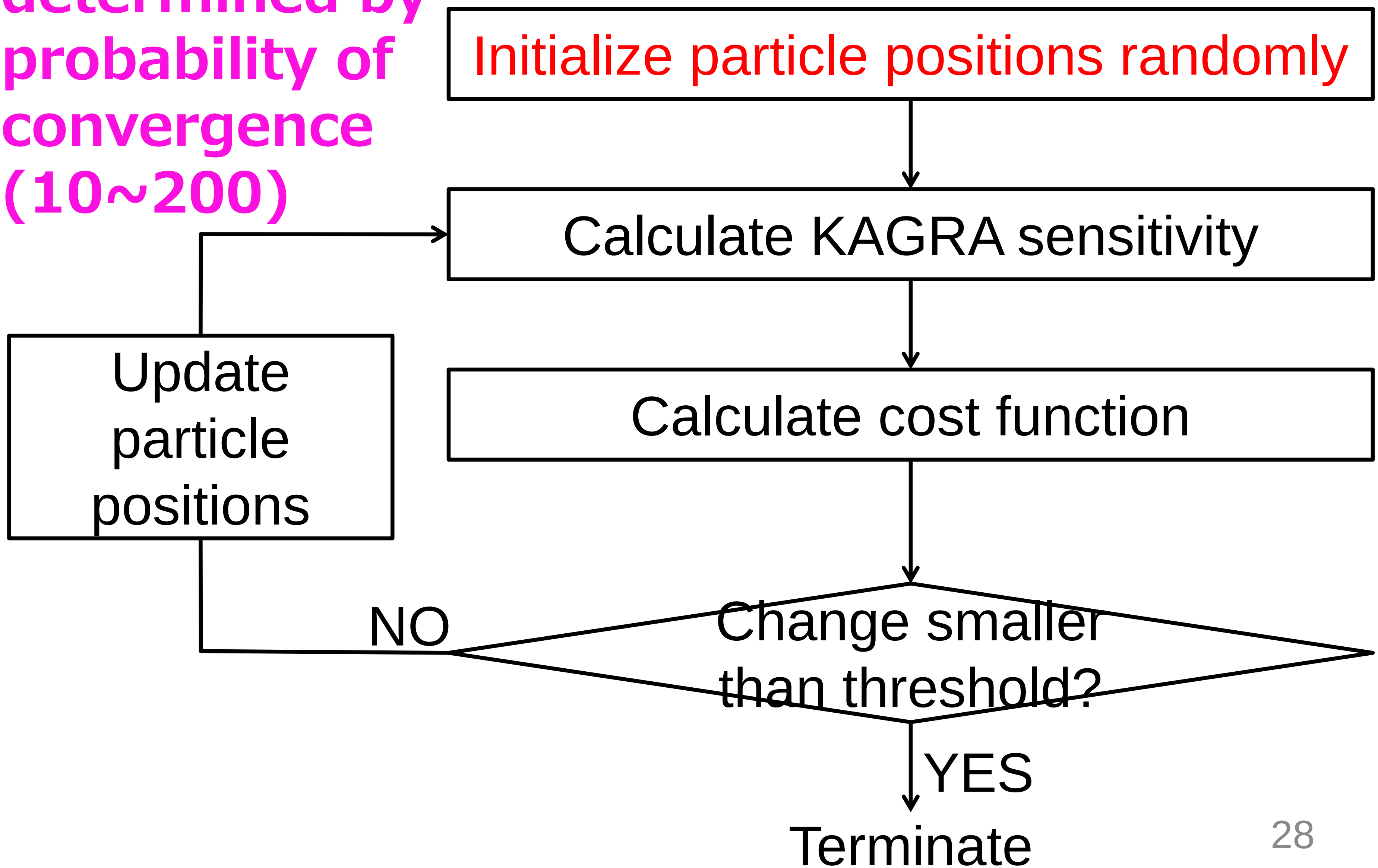
- Also applied to KAGRA+ study and showed optimized sensitivity for 3 candidates

  Sensitivity data available at [JGW-G1808426]

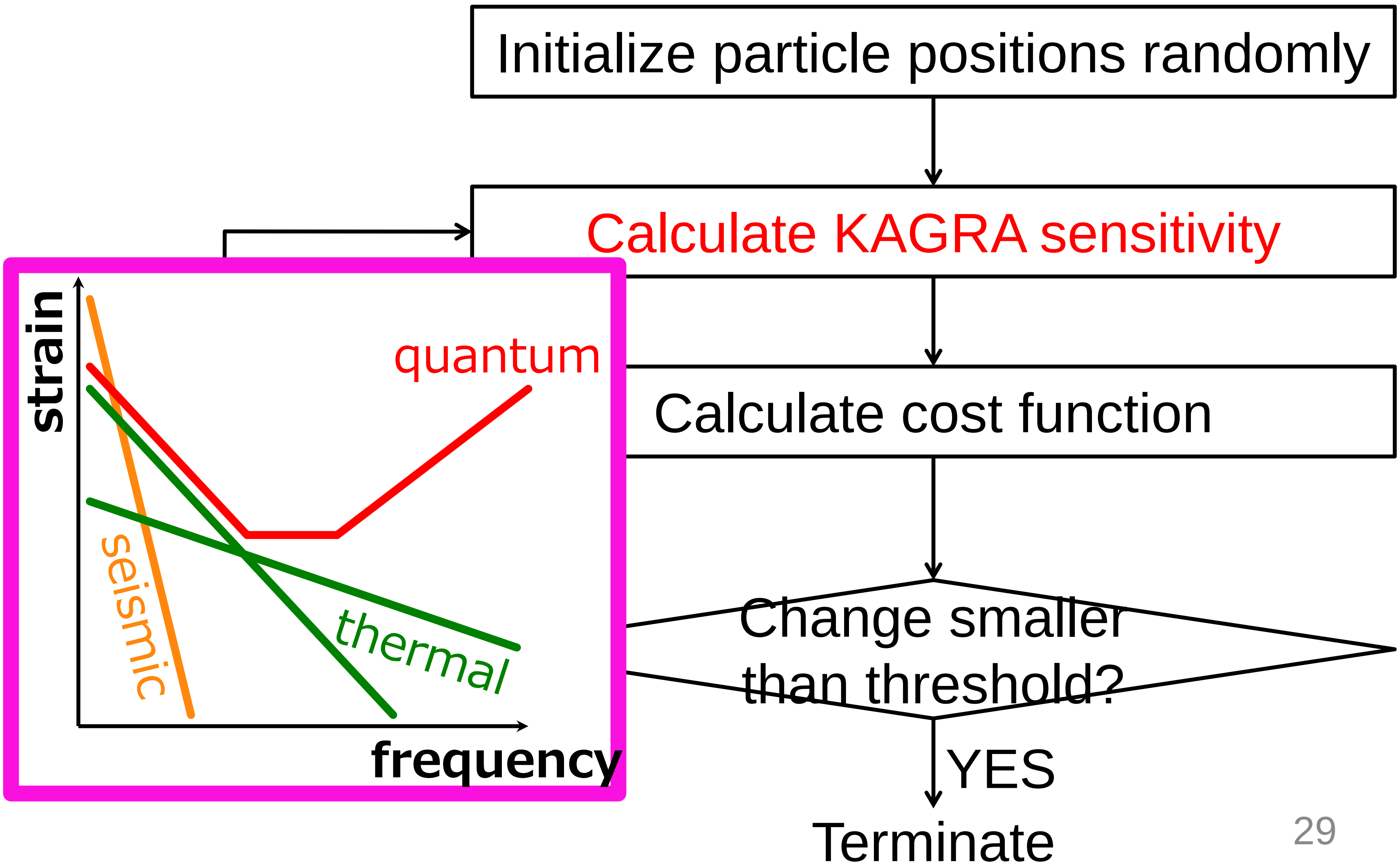
Supplementary Slides

PSO Algorithm

Swarm size determined by probability of convergence (10~200)



PSO Algorithm



PSO Algorithm

Initialize particle positions randomly

Calculate KAGRA sensitivity

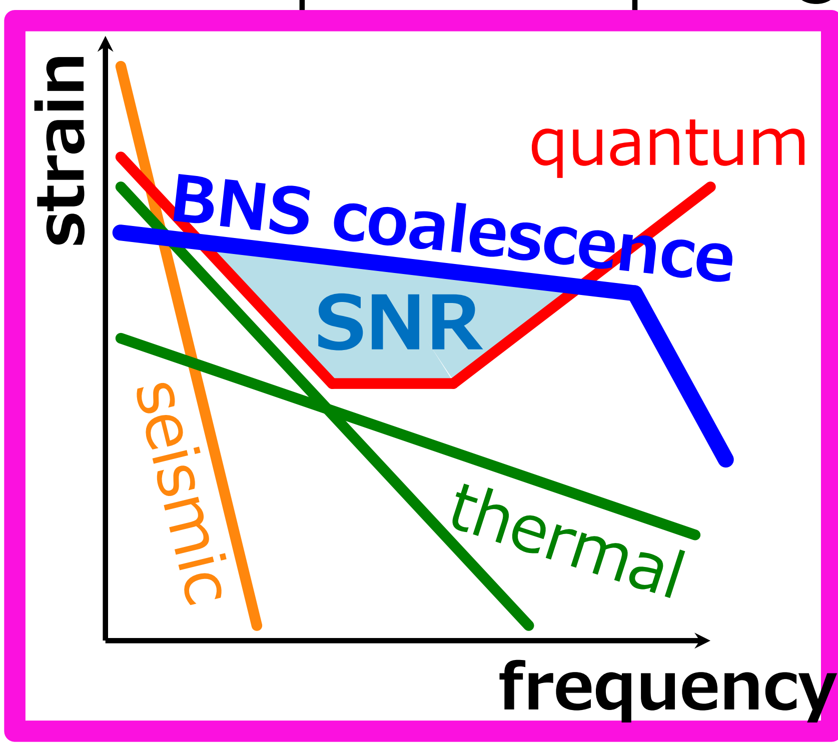
Calculate cost function

Change smaller than threshold?

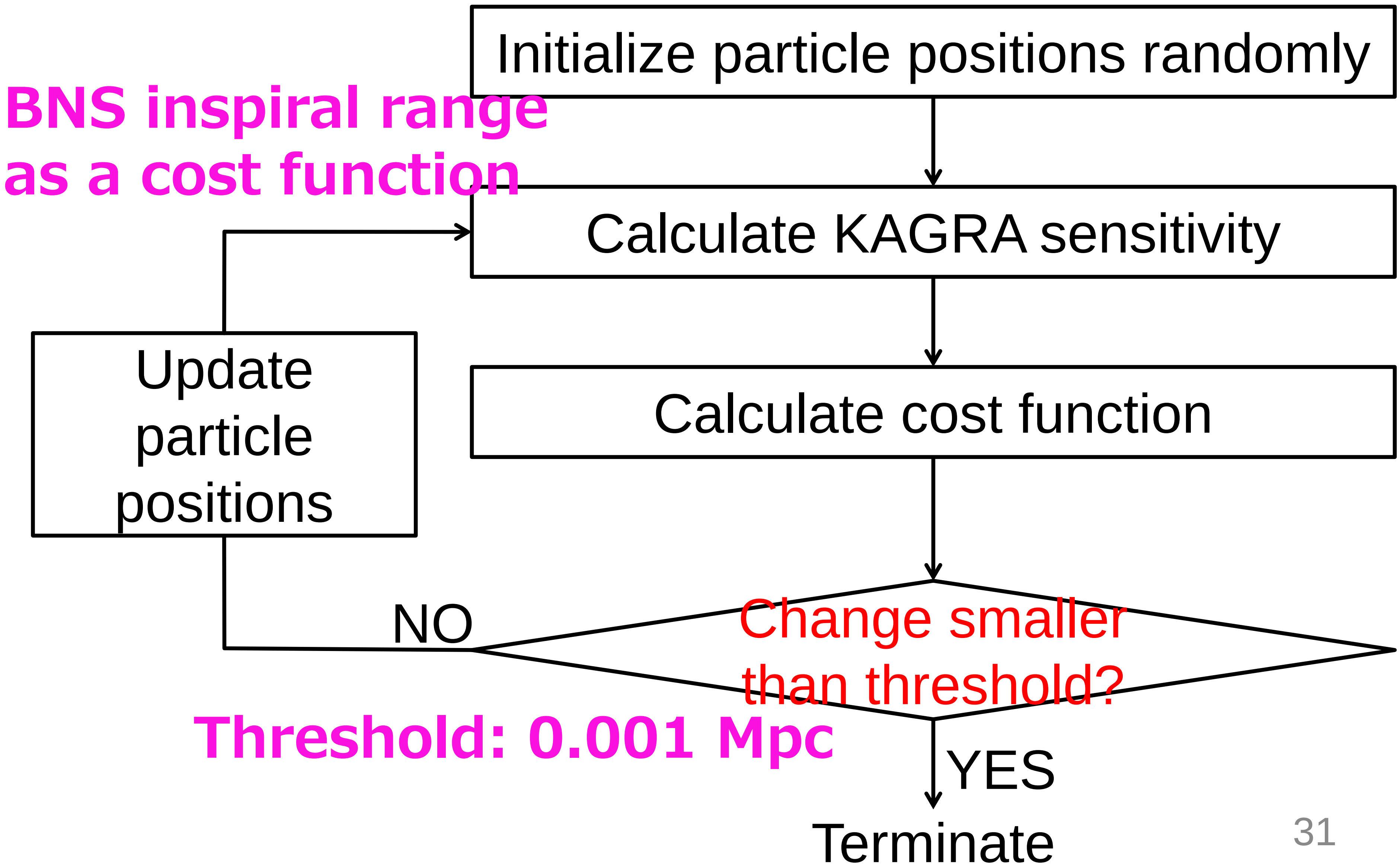
YES

Terminate

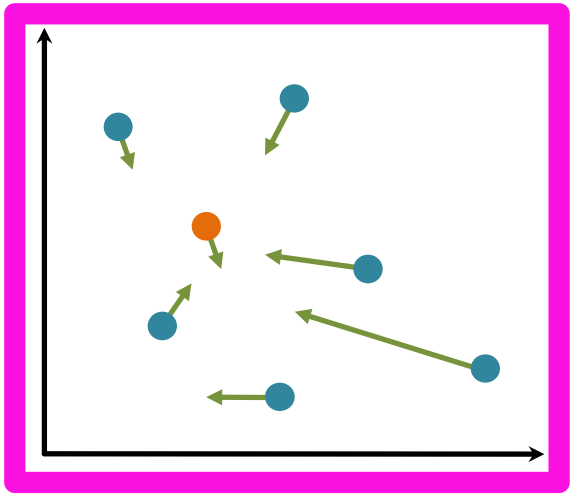
BNS inspiral range as a cost function



PSO Algorithm



PSO Algorithm



Initialize particle positions randomly

↓

Calculate KAGRA sensitivity

↓

Calculate cost function

↓

Change smaller than threshold?

NO → Update particle positions → Calculate KAGRA sensitivity

YES → Terminate

$$x_k(t+1) = x_k(t) + v_k(t)$$

$$v_k(t+1) = w v_k(t) + c_1 r_1 (\hat{x}_k - x_k(t)) + c_2 r_2 (\hat{x}_g - x_k(t))$$

Pyswarm

- Python package Pyswarm was used for this work
  https://pythonhosted.org/pyswarm/
  https://github.com/tisimst/pyswarm/
- PSO as easy as

```
xopt, fopt = pso(func, lb, ub)
```

- xopt: optimal parameter set
- fopt: optimal cost function
- func: cost function
- lb, ub: lower and upper bounds

PSO for GW Related Research

- CBC search

  Weerathunga & Mohanty, PRD 95, 124030 (2017)
  Wang & Mohanty, PRD 81, 063002 (2010)
  Bouffanais & Porter, PRD 93, 064020 (2016)

- Continuous GW search using pulsar timing array

  Wang, Mohanty & Jenet, ApJ 795, 96 (2014)

- Cosmological parameter estimation using CMB

  Prasad & Souradeep, PRD 85, 123008 (2012)

- Gravitational lens modeling

  Rogers & Fiege, ApJ 727, 80 (2011)

- Sensor correction filter design

  Conor Mow-Lowry, LIGO-G1700841 LIGO-T1700541

- Voyager quantum noise optimization

  input power, arm finesse, SRM transmissivity, homodyne, filter cavity

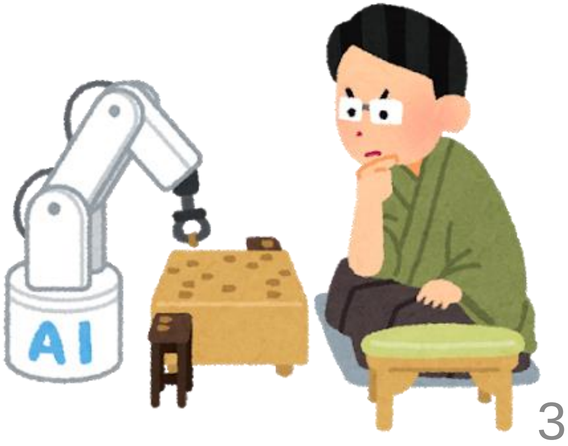
Pros and Cons of PSO

- Fast even for highly multidimensional parameter space ☺
  uses entire swarm's information to search
- Requires small number of design variables and little prior information ☺
  basically only swarm size and termination criterion
  prior information is only search range
- No guarantee for convergence to global maximum ☹
  stochastic method
- Do not give error of the parameters ☹
  no direct information on stability of the solution

→ Sounds great for detector design

Other Optimization Methods

- Simulated annealing
  tuning cooling schedule is troublesome
- Genetic algorithm
  too many design variables
- Markov chain Monte Carlo
  tend to be dependent on prior distribution
  gives error from posterior distribution
  takes time
- Machine learning
  if the problem well-modeled,
  you don’t need ML



Swarm Size Determination

- Probability of convergence: ratio of PSO trials resulted within 0.1 Mpc or 10^{-3} deg^2
- Increased swarm size until probability of convergence is larger than 90%

number of params	3	5	7
number of particles	10	20	200
number of iterations	52±13	73±16	60±18
probability of convergence	98 %	96 %	91 %

* From 100 PSO trials

IFO Parameter Search Range

	Lower bound	Upper bound	KAGRA Default	Precision
Detuning angle [deg]	86.5 (or 60) *	90	86.5	0.1
Homodyne angle [deg]	90	180	135.1	3
Mirror temperature [K]	20	30	22	0.09
Power attenuation	0.01	1	1	0.02
SRM reflectivity	0.5	1	0.92 (85%)	6e-4
Wire length [cm]	20	100	35	0.02
Wire safety factor	3	30	12.57 (0.8 mm)	0.07

* Considering SRC nonlinearity, maximum detuning is 3.5 deg (see Y. Aso+ CQG 29, 124008)

step size which changes BNS inspiral range by 0.1 Mpc

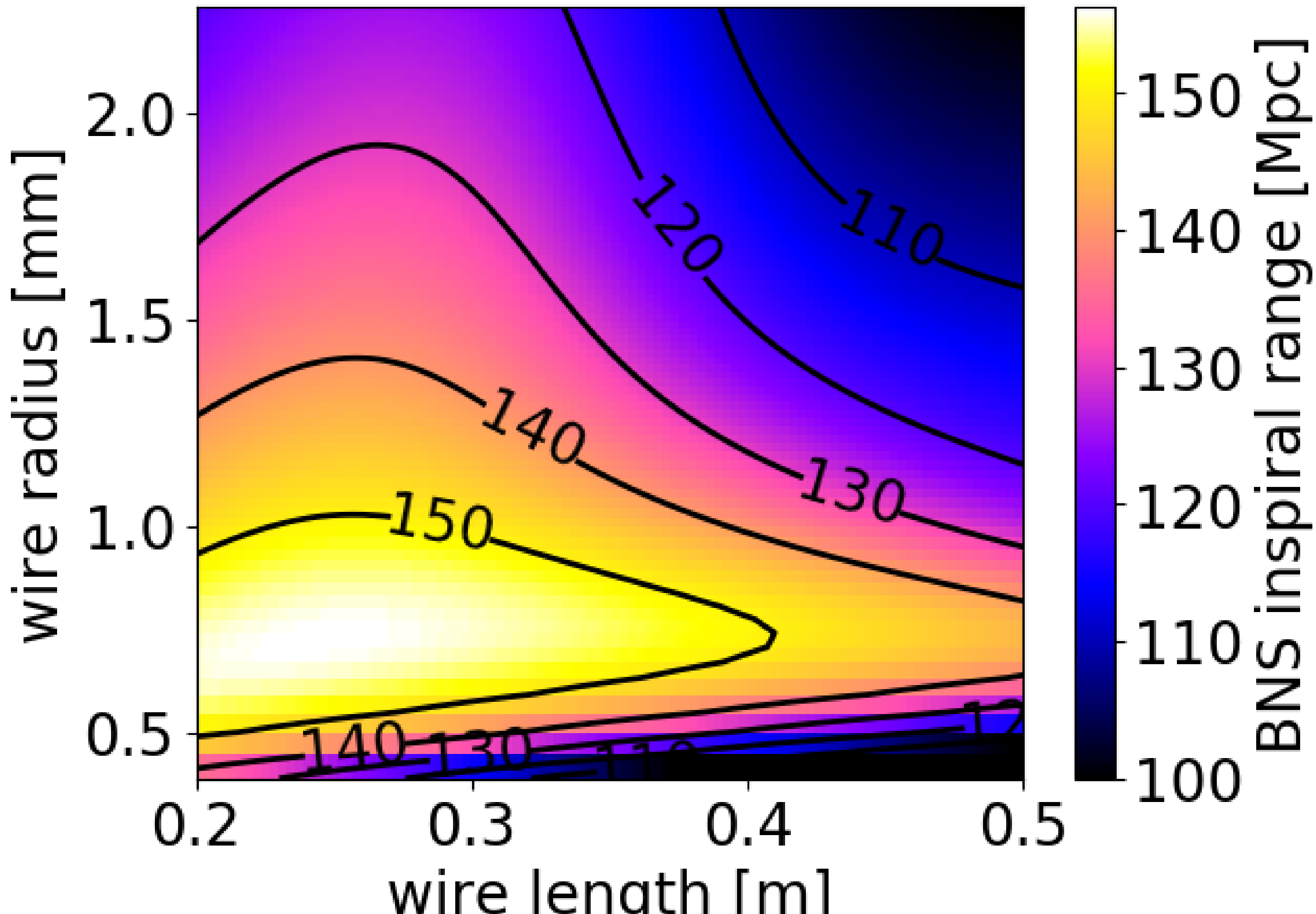
- *Reflecting wall* boundary:
  x=xmax, v=-v if x>xmax
  x=xmin, v=-v if x<xmin

Money

- Detuning angle and homodyne angle can be retuned without additional cost
- Mirror temperature and input power can be retuned without additional cost if power at BS is less than ~1 kW (~100 W entering PRM)
- Change in SRM reflectivity require ~0.1 Million USD
- Change in wire parameters require ~0.01 Million USD/fiber
- Change in wire length additionally require test mass suspension design change at ~0.1 Million USD/mirror
- Change in the test mass require ~0.6 Million USD/mirror (more for heavier ones)

Fiber Length and Diameter

- 25cm/φ1.4mm is optimum for BNS IR if other parameters are fixed (default: 35cm/φ1.6mm)



Fisher Matrix Analysis

- Fisher matrix

$$\Gamma_{ij} = 4\Re \int_{f_{\min}}^{f_{\max}} \sum_k \frac{\partial h_k^*(f)}{\partial \lambda^i} \frac{\partial h_k(f)}{\partial \lambda^j} \frac{\mathrm{d}f}{S_{\mathrm{n,k}}(f)}$$

- Covariance

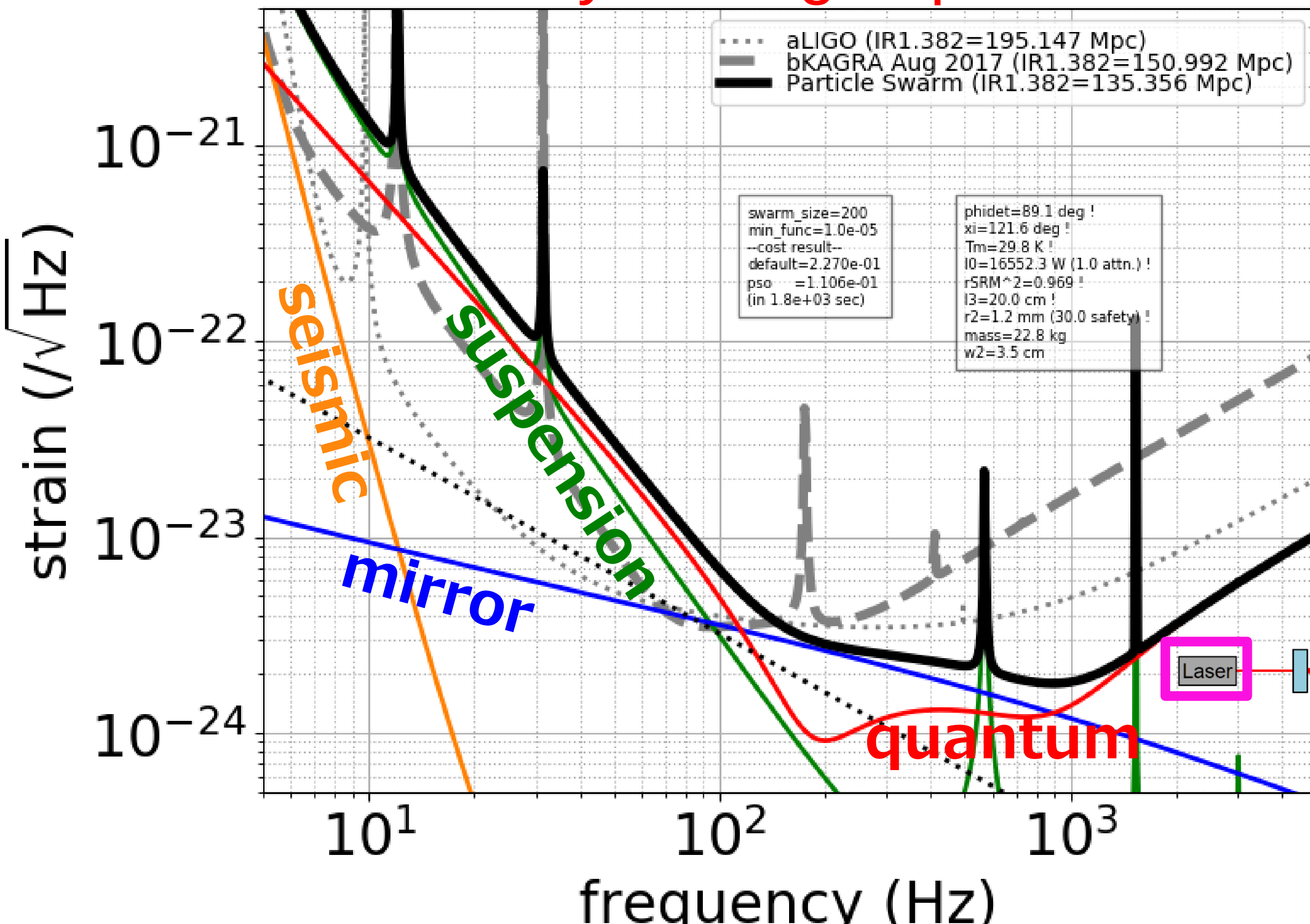
$$\sqrt{\langle(\delta\lambda^i \delta\lambda^j)\rangle} = \sqrt{(\Gamma^{-1})^{ij}}$$

- 11 binary parameters considered

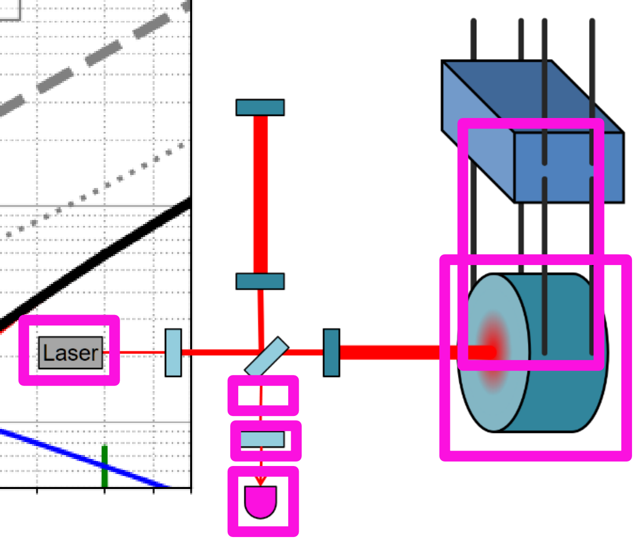
mc: chirp mass
eta: symmetric mass ratio
tc, phic: time and phase for coalescence
dL: luminosity distance
chis, chia: symmetric/asymmetric spin $\chi_{\mathrm{s/a}} = (\chi_1 \pm \chi_2)/2$
thetas, phis: colatitude / longitude of source
cthetai: inclination angle
psip: polarization angle

Optimization for Fixed Sky Location

- Result for fixed sky location and polarization angle is similar to sky average optimization



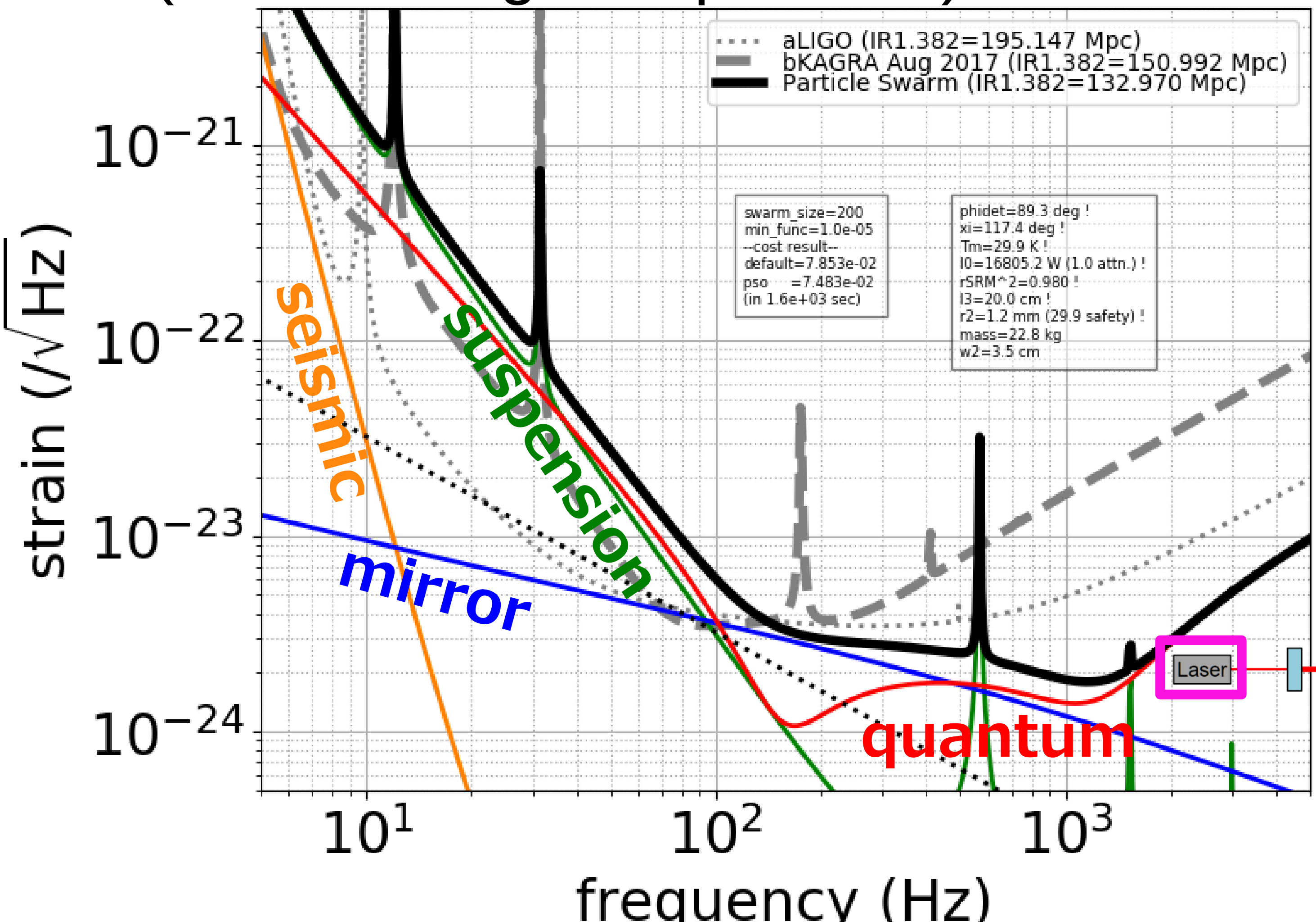
shortest and thickest fibers
1.7 kW at BS
30 K (max)



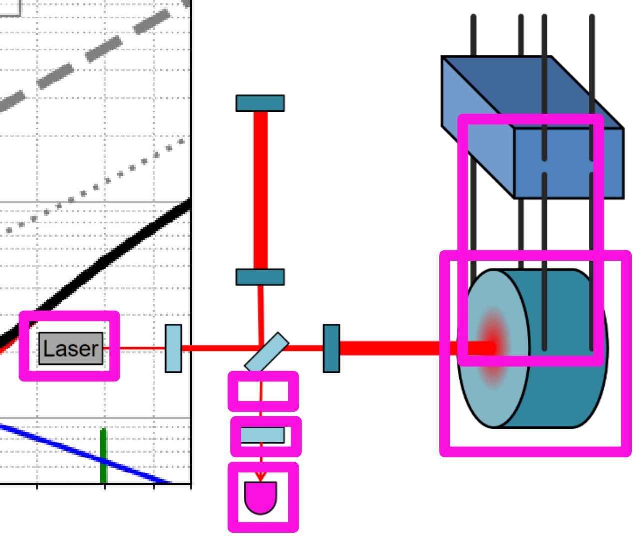
Symmetric Spin Optimization

- Similar to sky localization optimization (focus on high frequencies)

$$\chi_s = \frac{\chi_1 + \chi_2}{2}$$



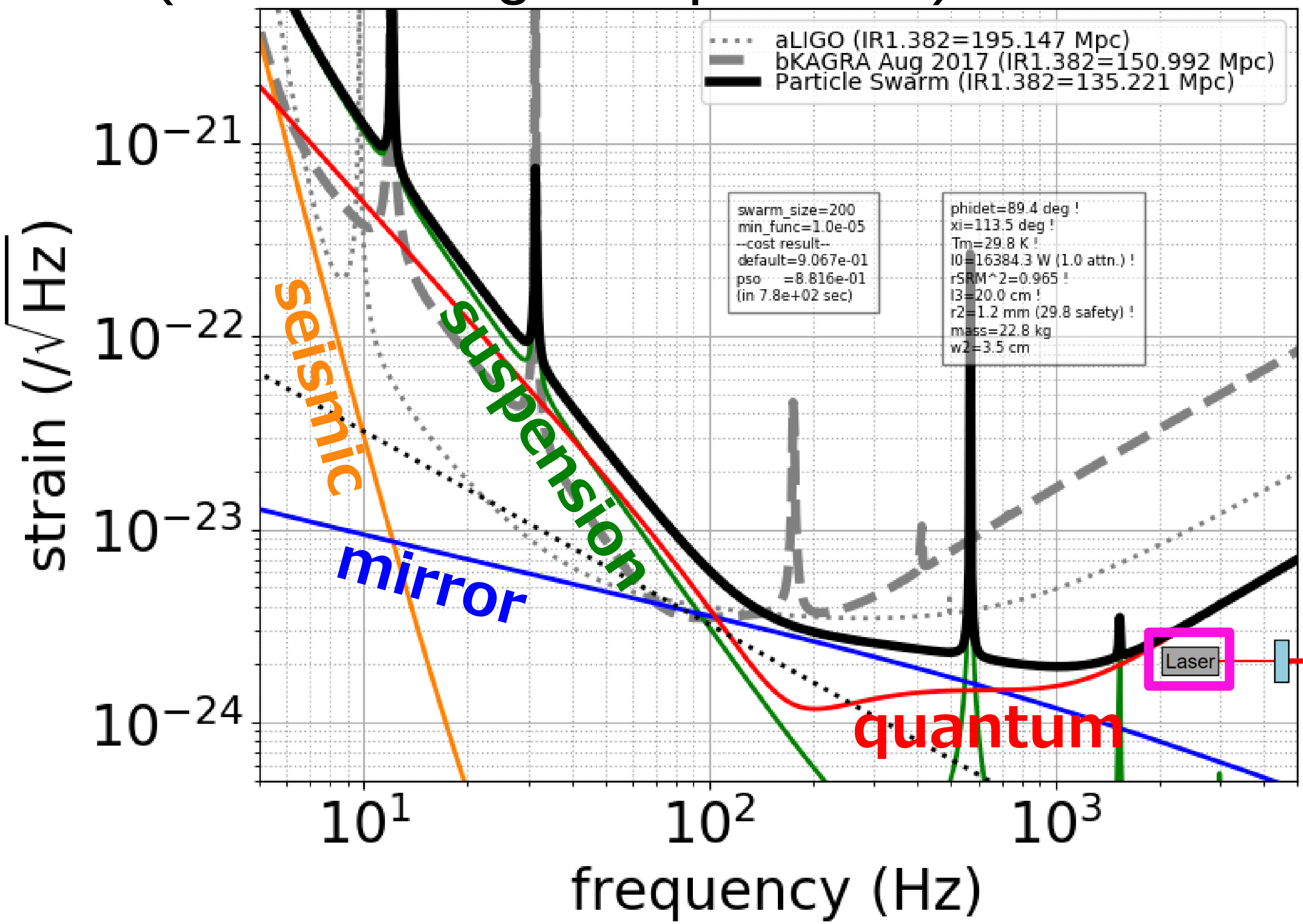
shortest and thickest fibers
1.7 kW at BS
30 K (max)



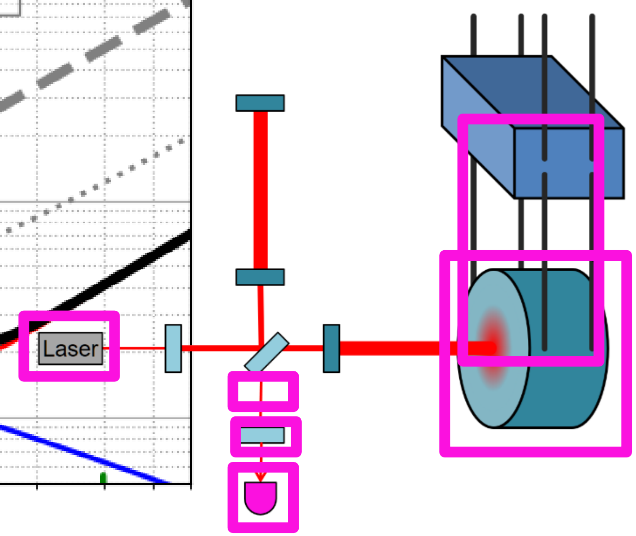
Asymmetric Spin Optimization

- Similar to sky localization optimization (focus on high frequencies)

$$\chi_a = \frac{\chi_1 - \chi_2}{2}$$

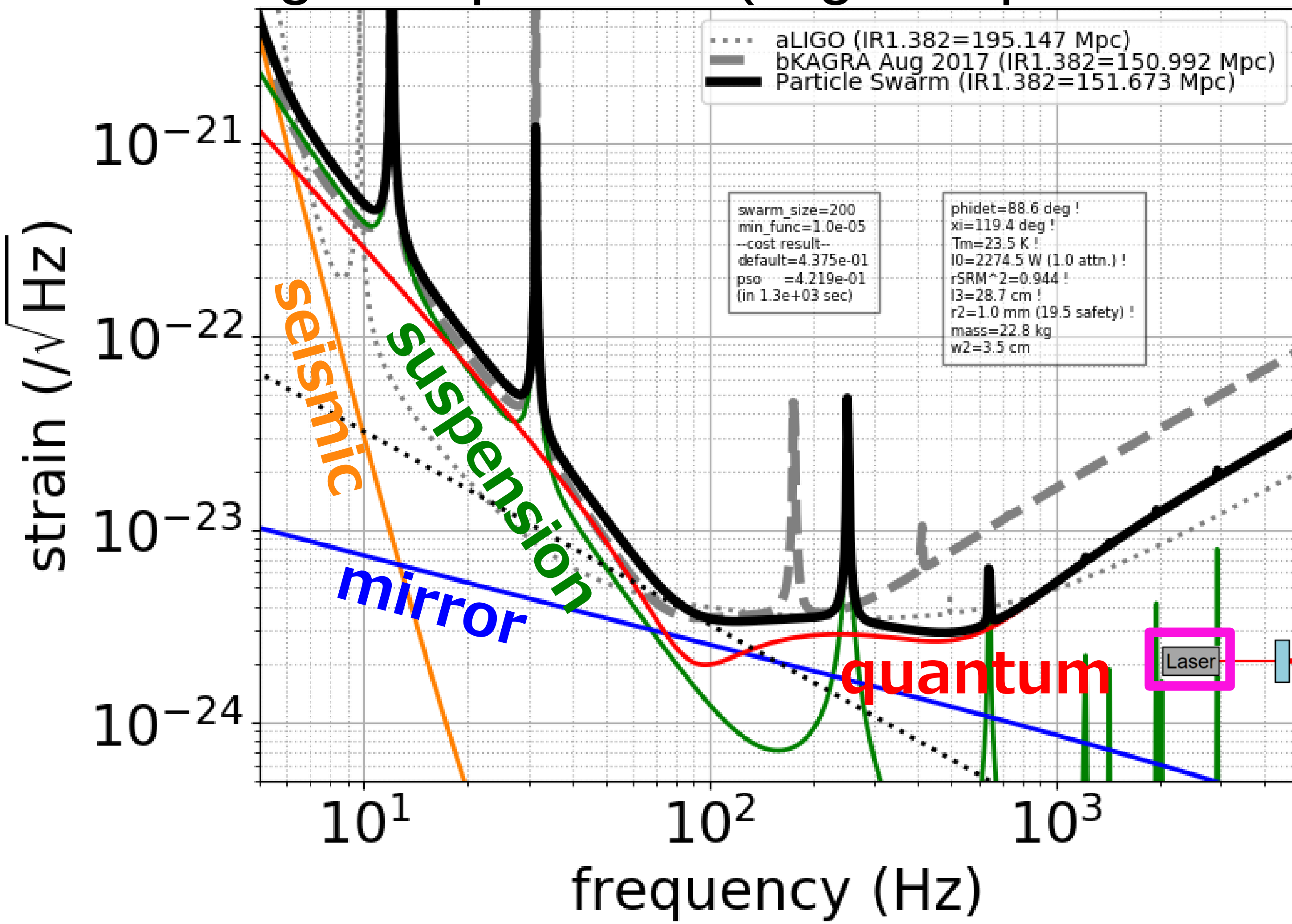


shortest and thickest fibers
1.7 kW at BS
30 K (max)

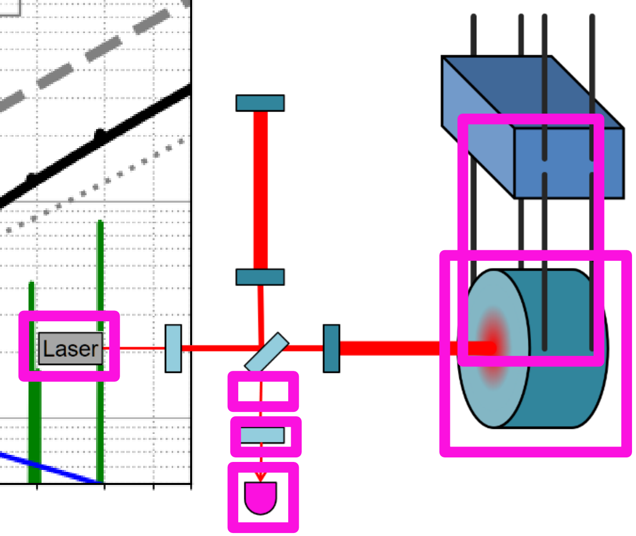


Distance Optimization

- Similar to inspiral range optimization, but slight shift to high frequencies (slight improvement by 2%)

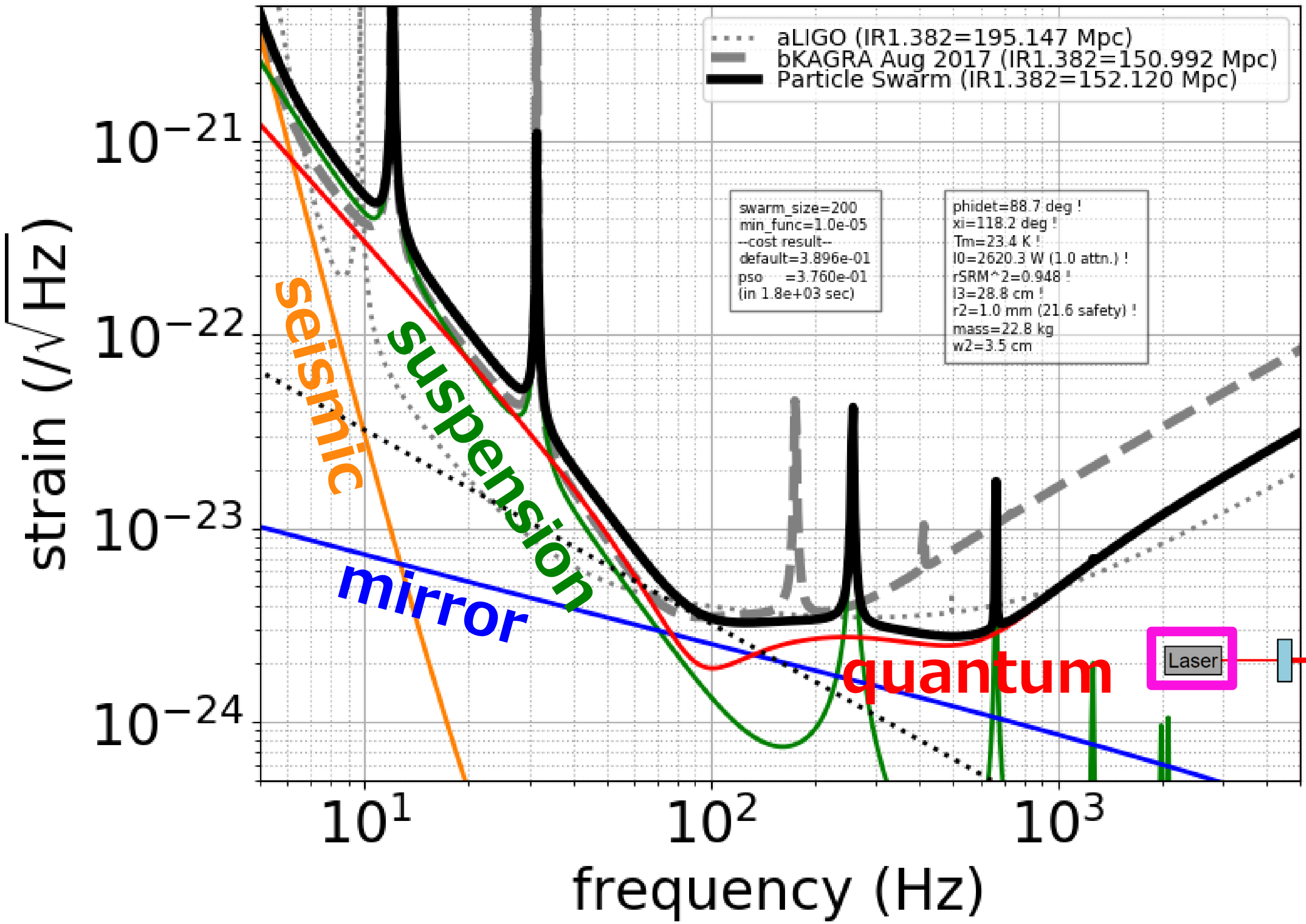


shorter and thicker fibers
2.3 kW at BS
24 K

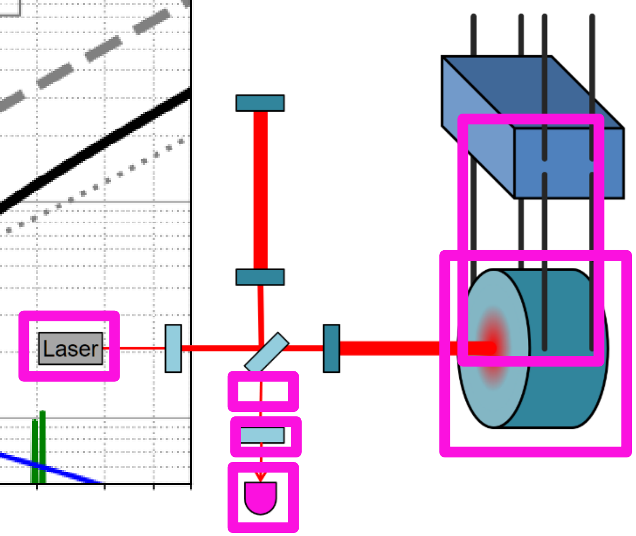


Inclination Angle Optimization

- Similar to distance optimization (PE degeneracy)

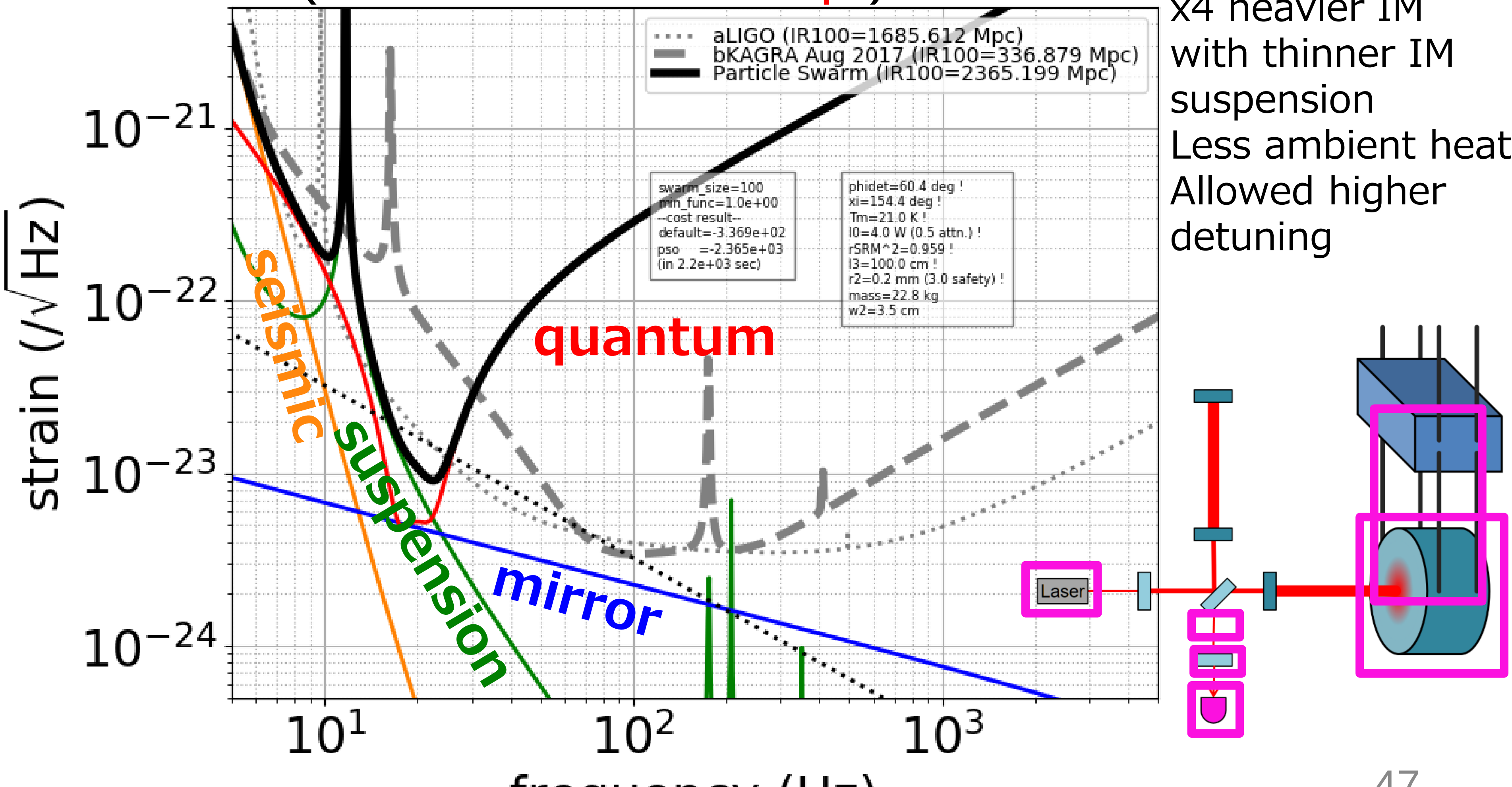


shorter and thicker fibers
2.3 kW at BS
24 K



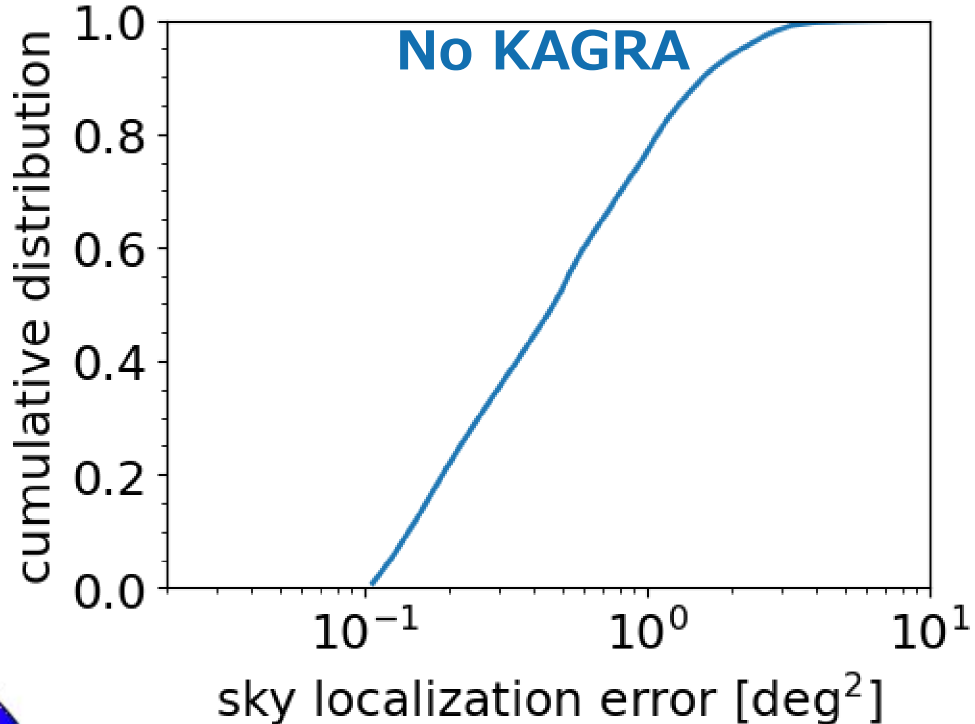
BBH100 IR Optimization

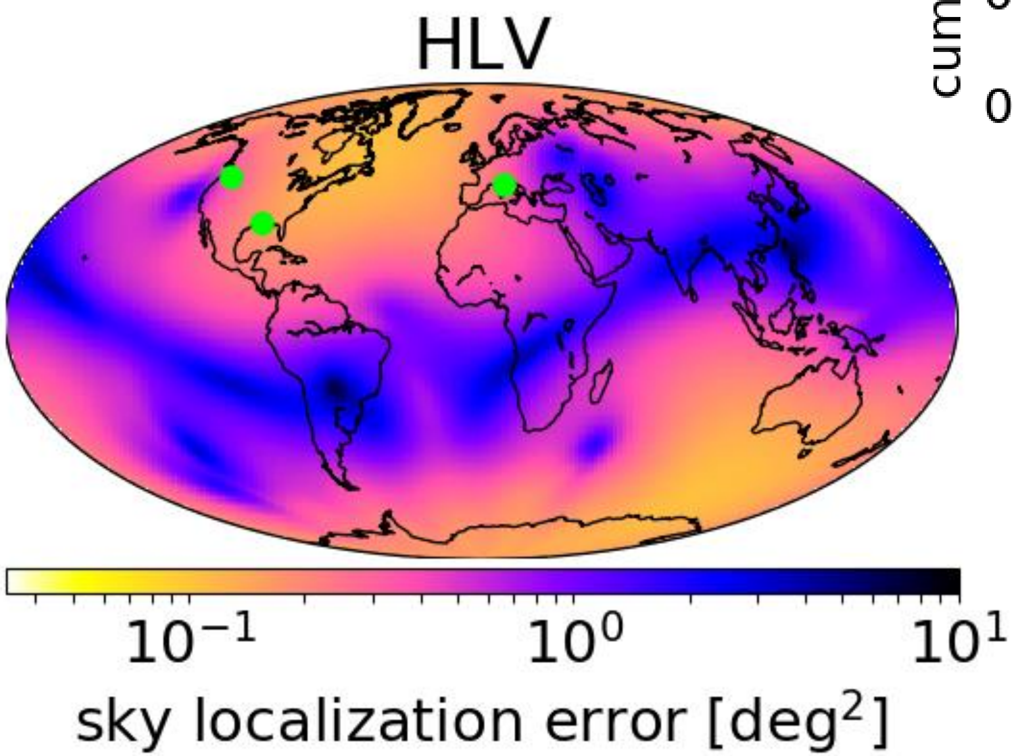
- Low power, low temperature with thin and longer fibers (KAGRA+ LF concept)



x4 heavier IM with thinner IM suspension
Less ambient heat
Allowed higher detuning

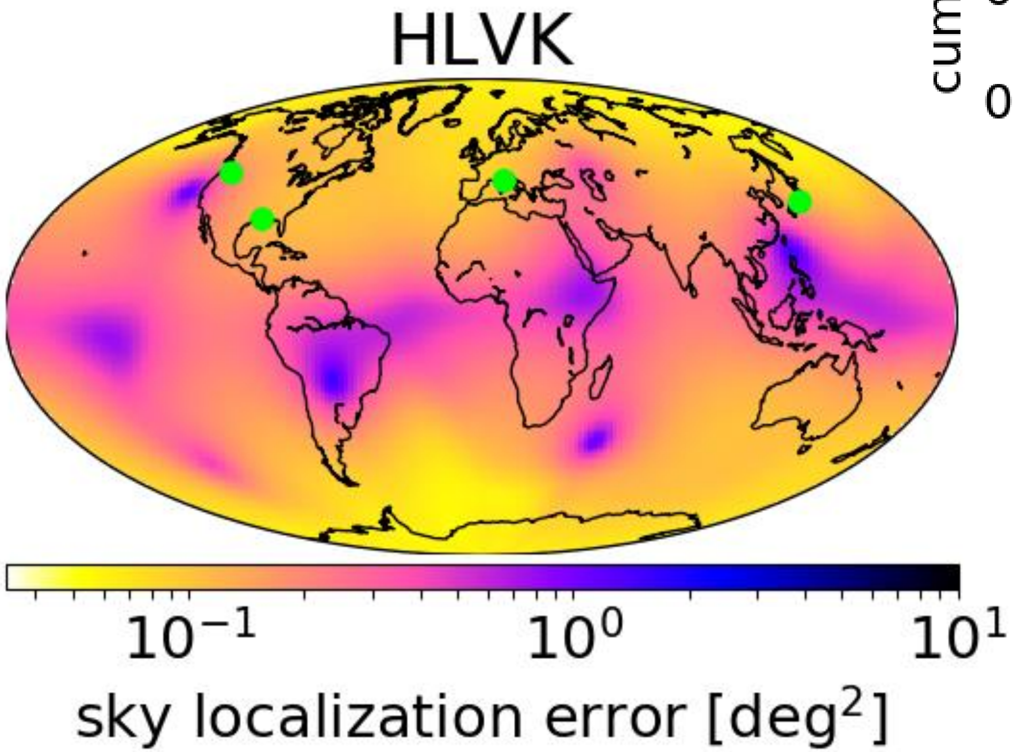
Sky Localization with HLV

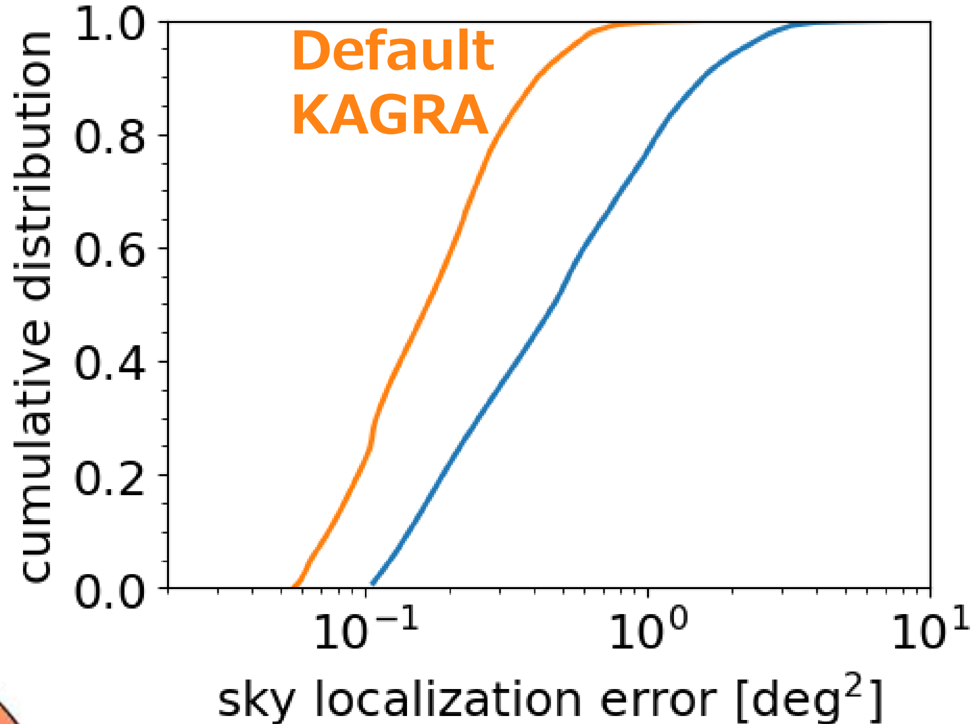




	median
HLV	0.472 deg^2
HLVK	
HLVK+	

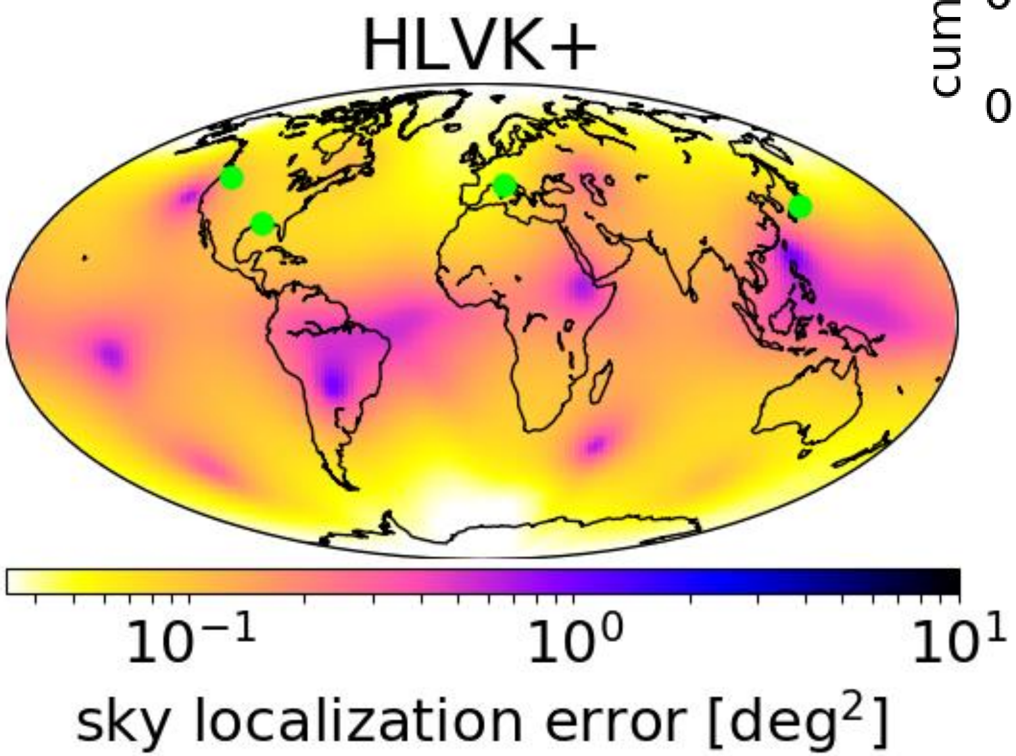
Sky Localization with HLVK

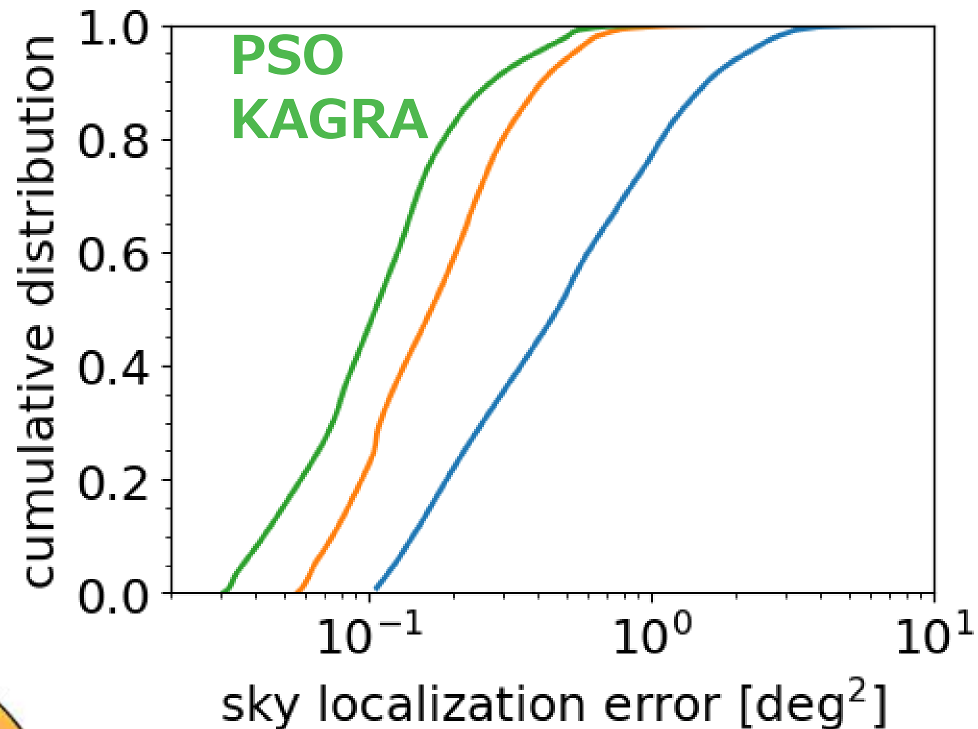




	median
HLV	0.472 deg^2
HLVK	0.168 deg^2
HLVK+	

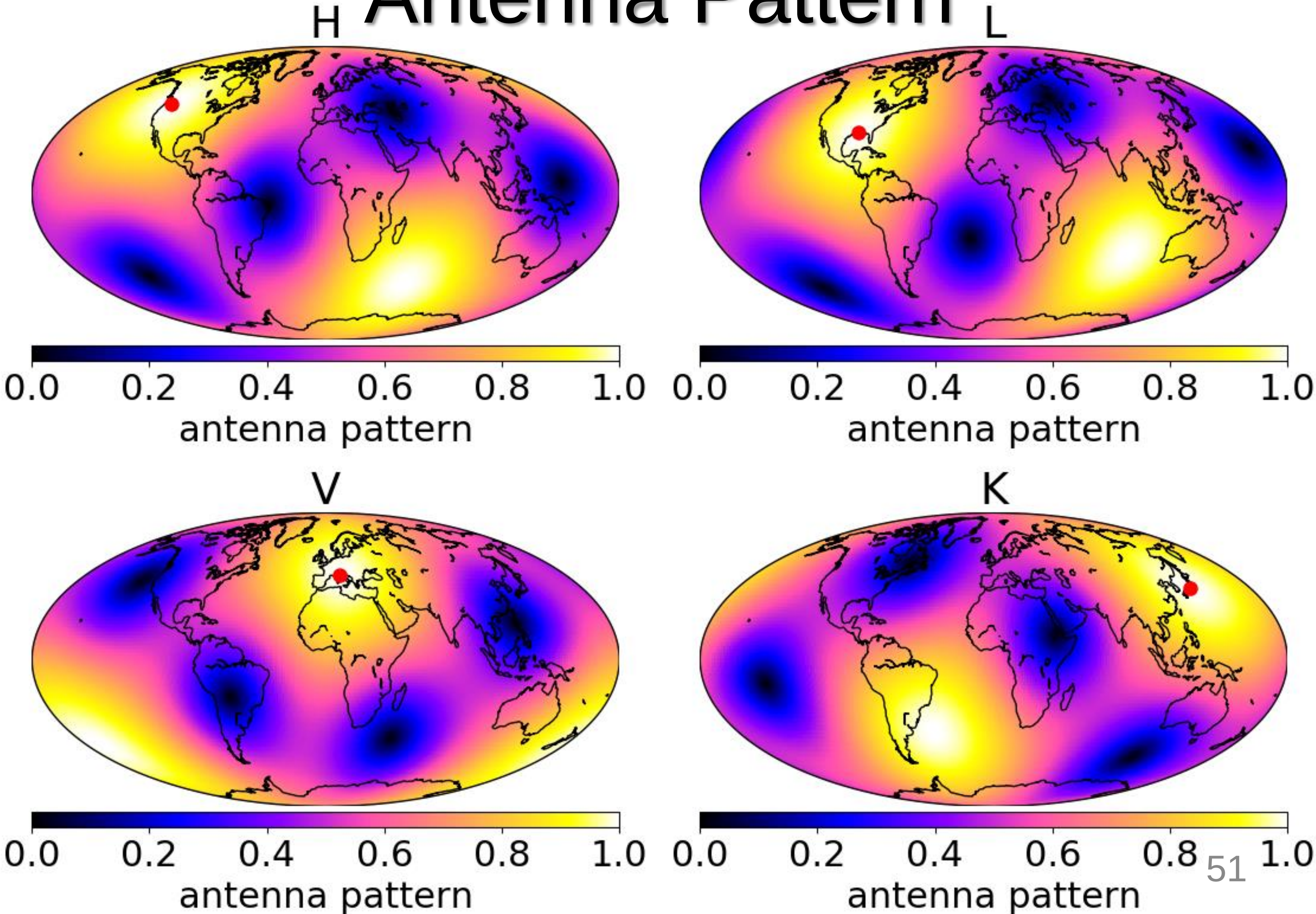
Sky Localization with HLVK+





	median
HLV	0.472 deg^2
HLVK	0.168 deg^2
HLVK+	0.107 deg^2

Antenna Pattern



2G/2G+ Parameter Comparison

	KAGRA	AdVirgo	aLIGO	A+	Voyager
Arm length [km]	3	3	4	4	4
Mirror mass [kg]	23	42	40	80	200
Mirror material	Sapphire	Silica	Silica	Silica	Silicon
Mirror temp [K]	22	295	295	295	123
Sus fiber	35cm Sap.	70cm SiO_2	60cm SiO_2	60cm SiO_2	60cm Si
Fiber type	Fiber	Fiber	Fiber	Fiber	Ribbon
Input power [W]	67	125	125	125	140
Arm power [kW]	340	700	710	1150	3000
Wavelength [nm]	1064	1064	1064	1064	2000
Beam size [cm]	3.5 / 3.5	4.9 / 5.8	5.5 / 6.2	5.5 / 6.2	5.8 / 6.2
SQZ factor	0	0	0	6	8
F. C. length [m]	none	none	none	16	300

LIGO parameters from LIGO-T1600119, AdVirgo parameters from JPCS 610, 01201 (2015)

KAGRA Detailed Parameters

K. Komori *et al.*, JGW-T1707038

- **Optical parameters**
  - Mirror transmission: 0.4 % for ITM, 10 % for PRM, 15.36 % for SRM
  - Power at BS: 674 W
  - Detune phase: 3.5 deg (DRSE case)
  - Homodyne phase: 135.1 deg (DRSE case)
- **Sapphire mirror parameters**
  - TM size: 220 mm dia., 150 mm thick
  - TM mass: 22.8 kg
  - TM temperature: 22 K
  - Beam radius at ITM: 3.5 cm
  - Beam radius at ETM: 3.5 cm
  - Q of mirror substrate: 1e8
  - Coating: tantala/silica
  - Coating loss angle: 3e-4 for silica, 5e-4 for tantala
  - Number of layers: 22 for ITM, 40 for ETM
  - Coating absorption: 0.5 ppm
  - Substrate absorption: 50 ppm/cm
- **Suspension parameters**
  - TM-IM fiber: 35 cm long, 1.6 mm dia.
  - IM temperature: 16 K
  - Heat extraction: 5800 W/m/K at 20 K
  - Loss angle: 5e-6/2e-7/7e-7 for CuBe fiber/sapphire fiber/sapphire blade
- **Inspiral range calculation**
  - SNR=8, fmin=10 Hz, sky average constant 0.442478
- Seismic noise curve includes vertical coupling, vibration from heatlinks and Newtonian noise from surface and bulk

KAGRA Cryopayload

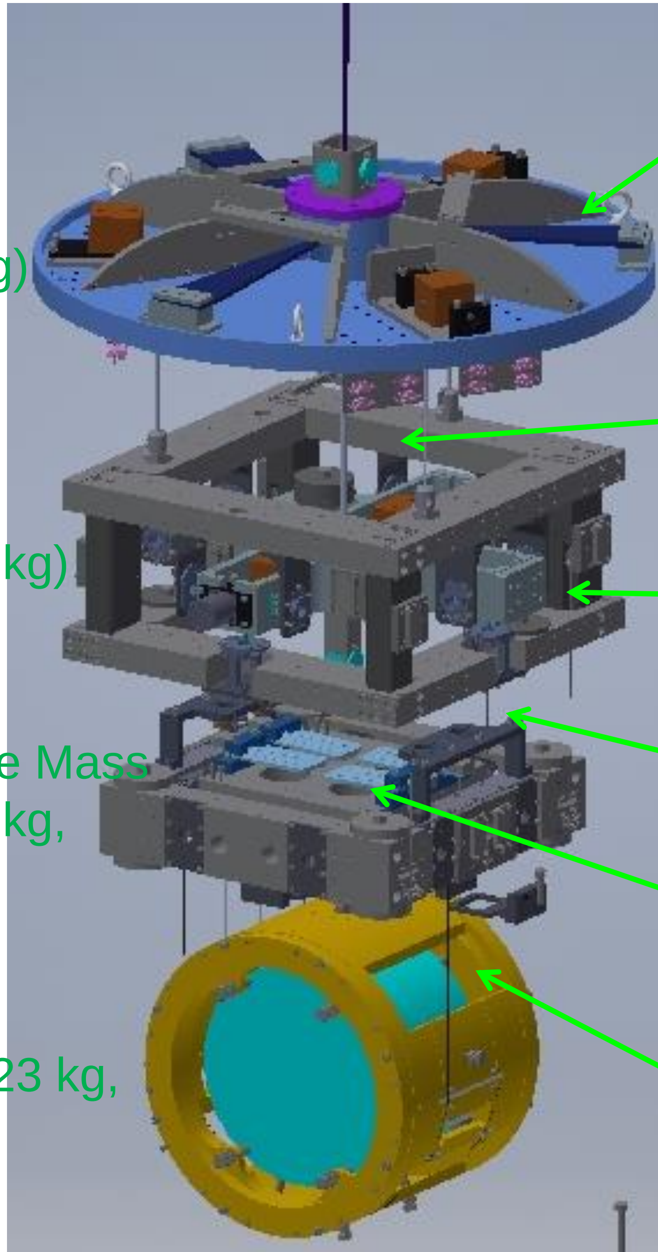
Provided by T. Ushiba and T. Miyamoto

Platform
(SUS, 65 kg)

Marionette
(SUS, 22.5 kg)

Intermediate Mass
(SUS, 20.1 kg,
16 K)

Test Mass
(Sapphire, 23 kg,
22 K)



3 CuBe blade springs

MN suspended by 1 Maraging steel fiber
(35 cm long, 2-7mm dia.)
MRM suspended by 3 CuBe fibers

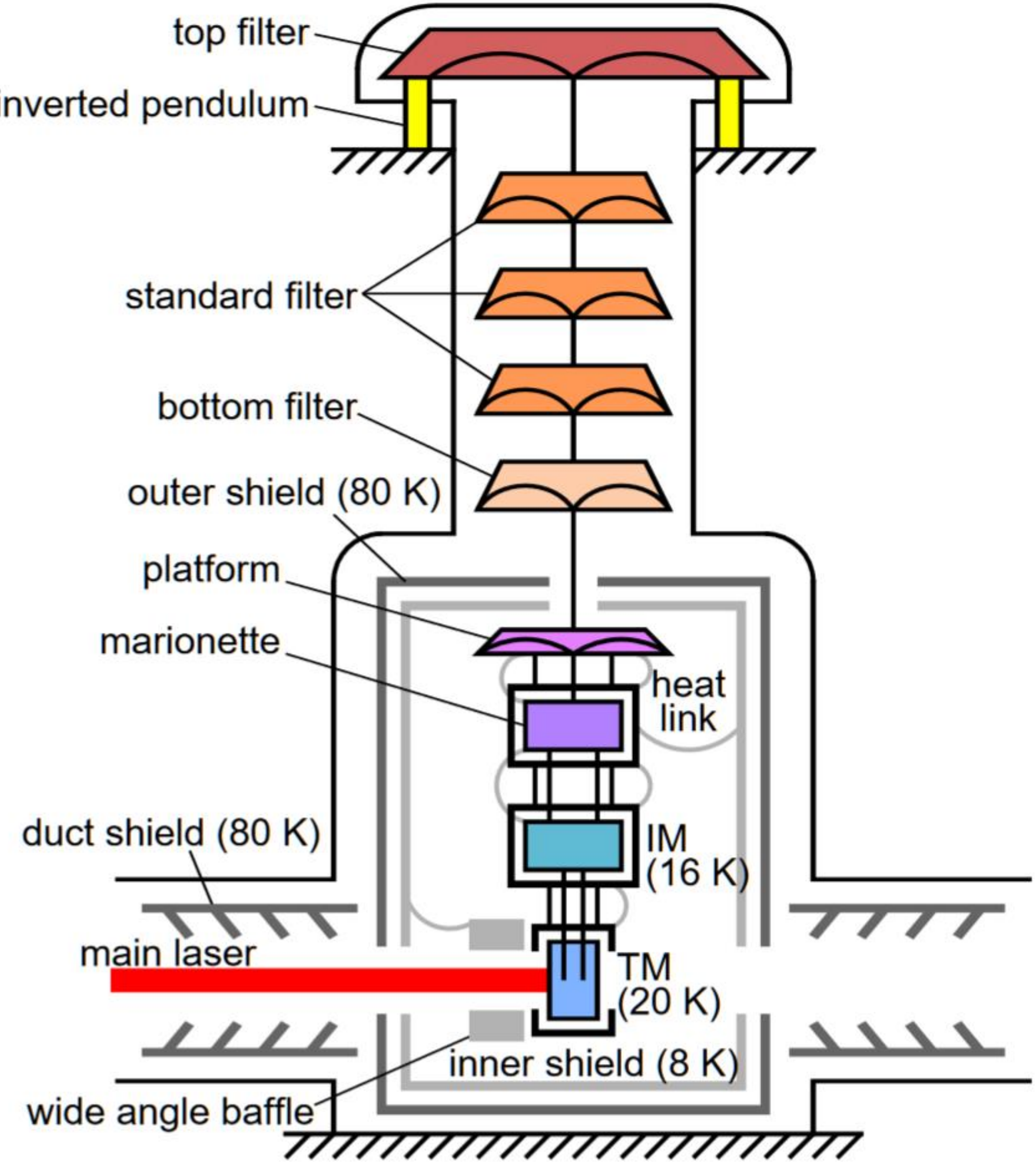
Heat link attached to MN

IM suspended by 4 CuBe fibers
(24 cm long, 0.6 mm dia)
IRM suspended by 4 CuBe fibers

4 sapphire blades

TM suspended by 4 sapphire fibers
(35 cm long, 1.6 mm dia.)
RM suspended by 4 CuBe fibers

KAGRA Cryostat Schematic



arXiv:1710.04823